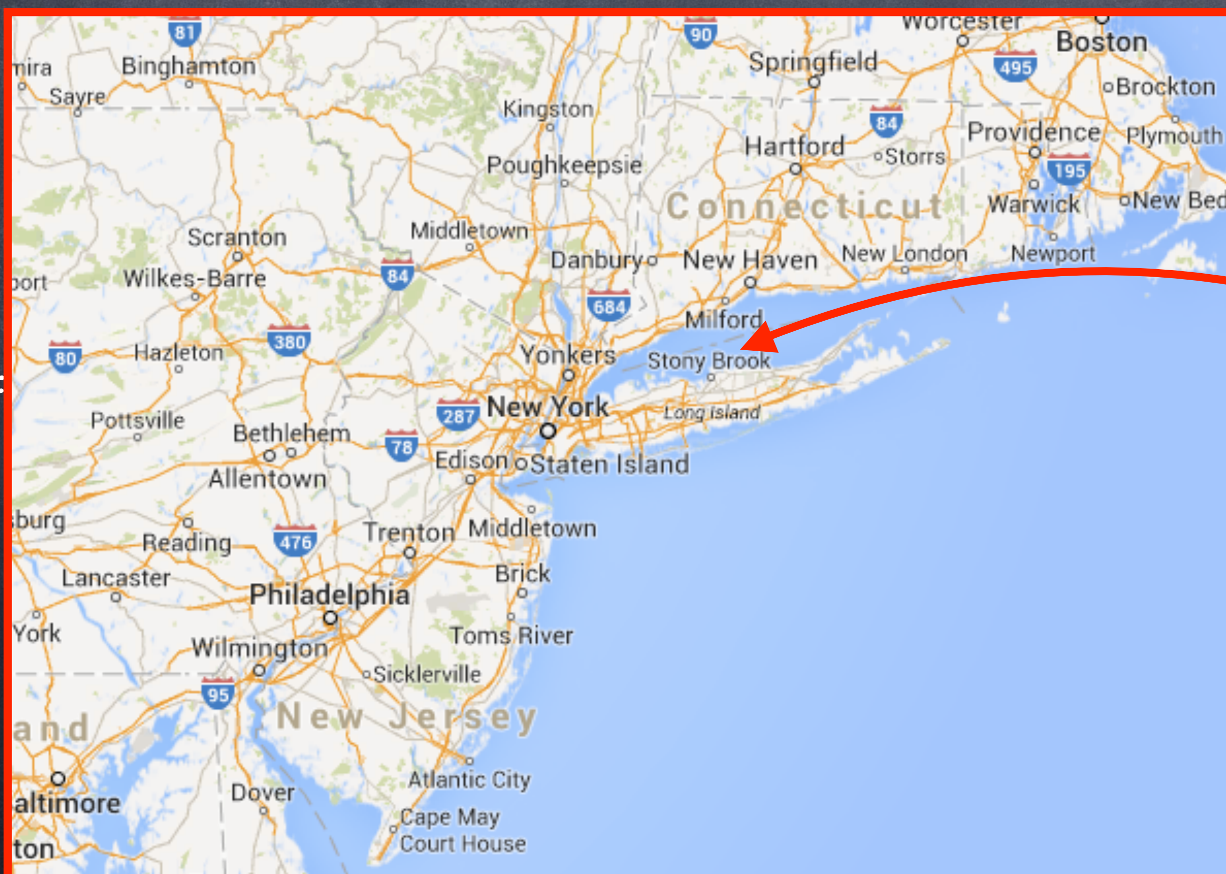
Neutrinos in Large-scale Structure

Marilena LoVerde
University of Chicago
(Fall 2015 —> C. N. Yang Institute for Theoretical Physics, Stony Brook University)

Neutrinos in Large-scale

(... Stony Brook University)

Neutrinos in Large-scale



(F

s, Stony Brook University)

Neutrinos in Large-scale structure

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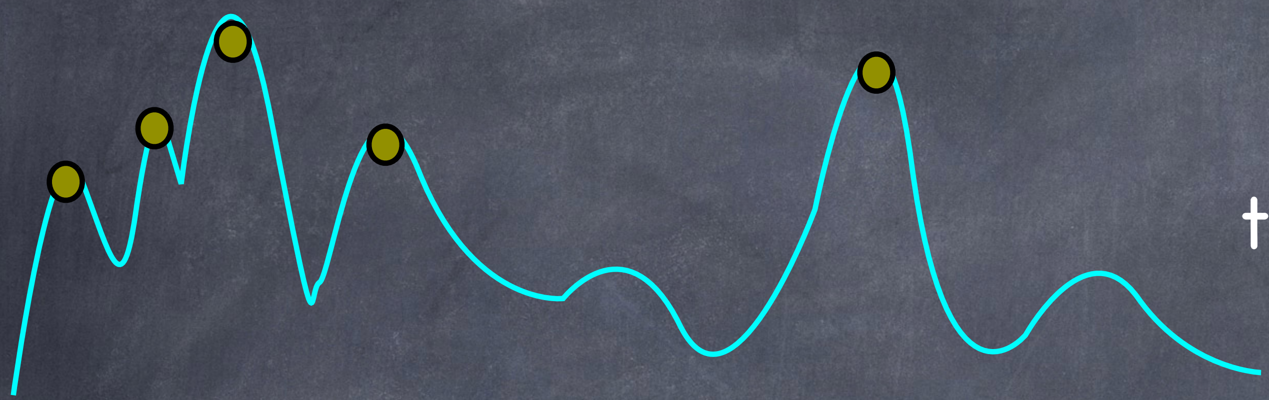
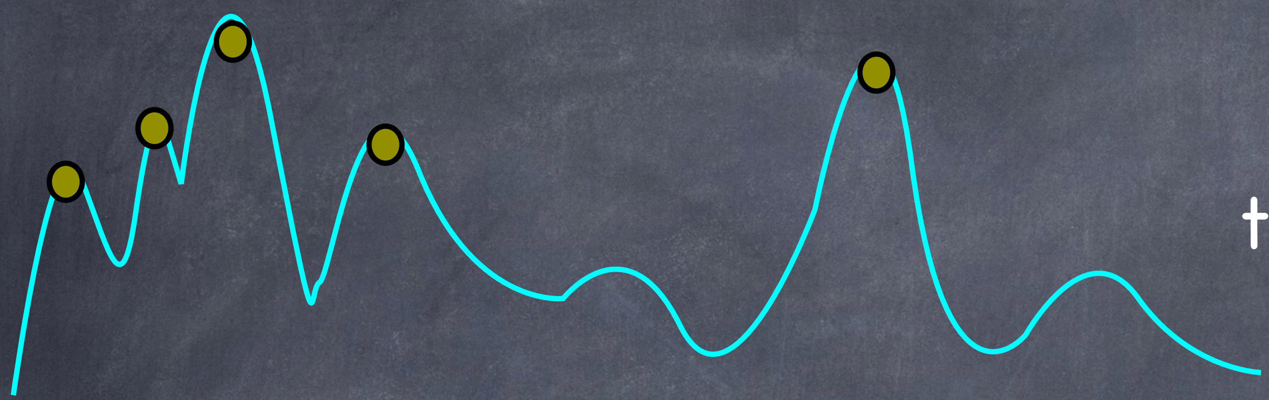
ML and Zaldarriaga 1310.6459
ML 1405.4855
ML 1404:4858
ML in prep.

Main idea:



halos are biased tracers of the matter density field

Main idea:



halos are biased tracers of the matter density field

the number density of halos is modulated by long-wavelength fluctuations in the matter density field

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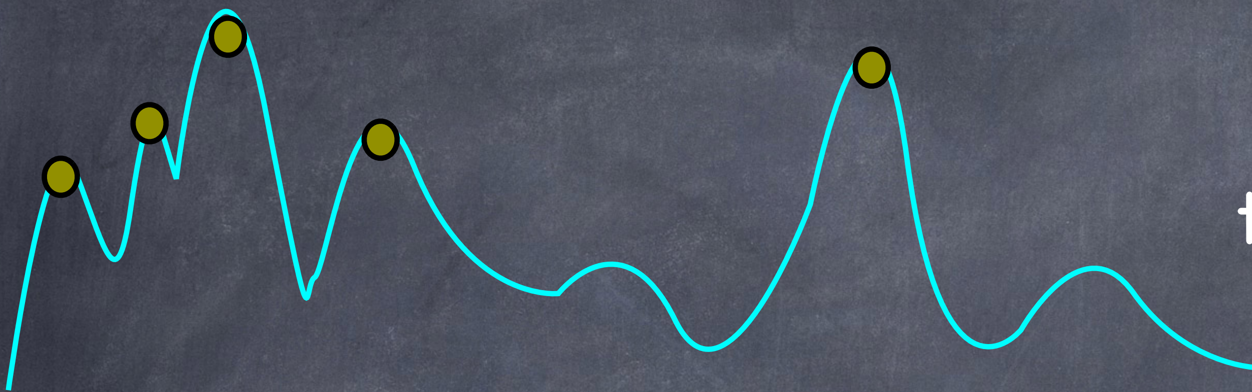
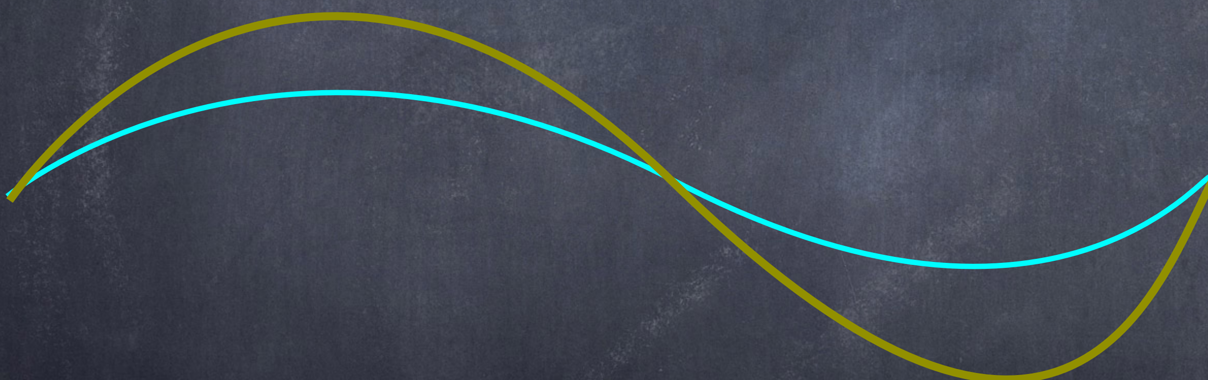


halos are biased tracers of the matter density field



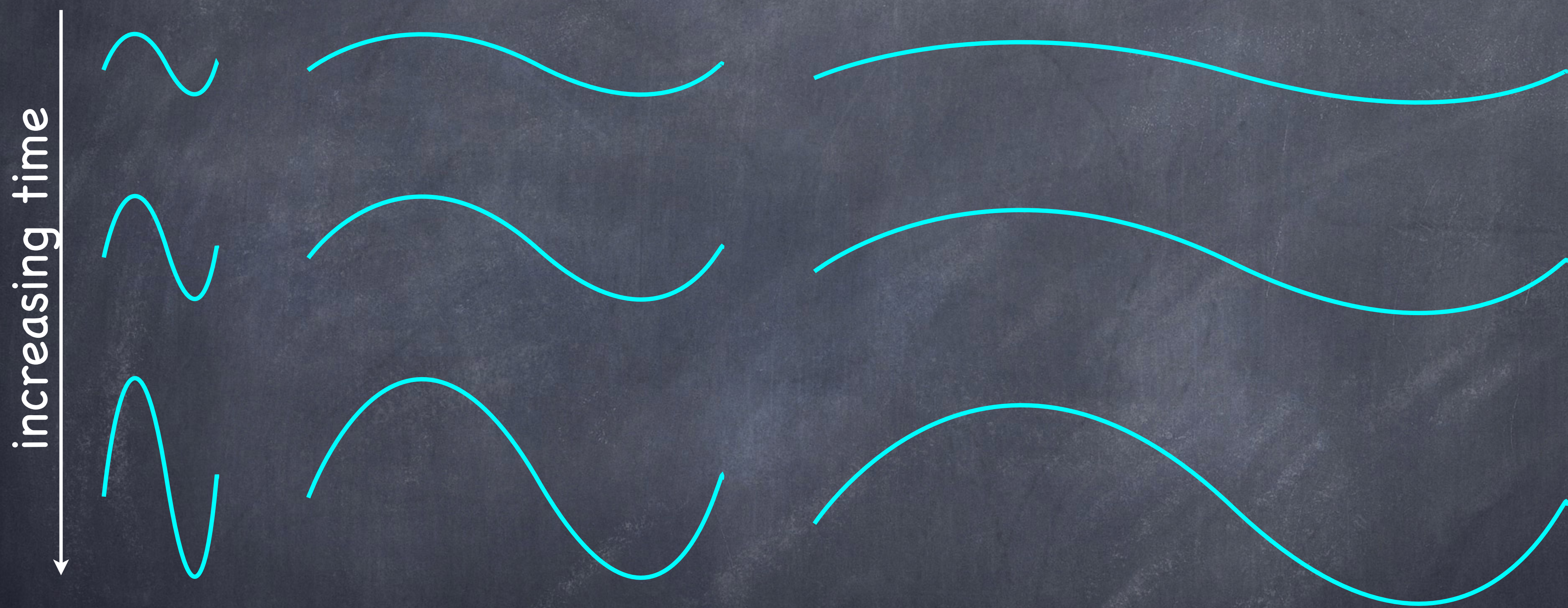
the number density of halos is modulated by long-wavelength fluctuations in the matter density field

$$\frac{\delta n}{n} \equiv b \left.\frac{\delta\rho}{\rho}\right|_{\text{long-wavelength}}$$

b is the halo bias

Main idea:

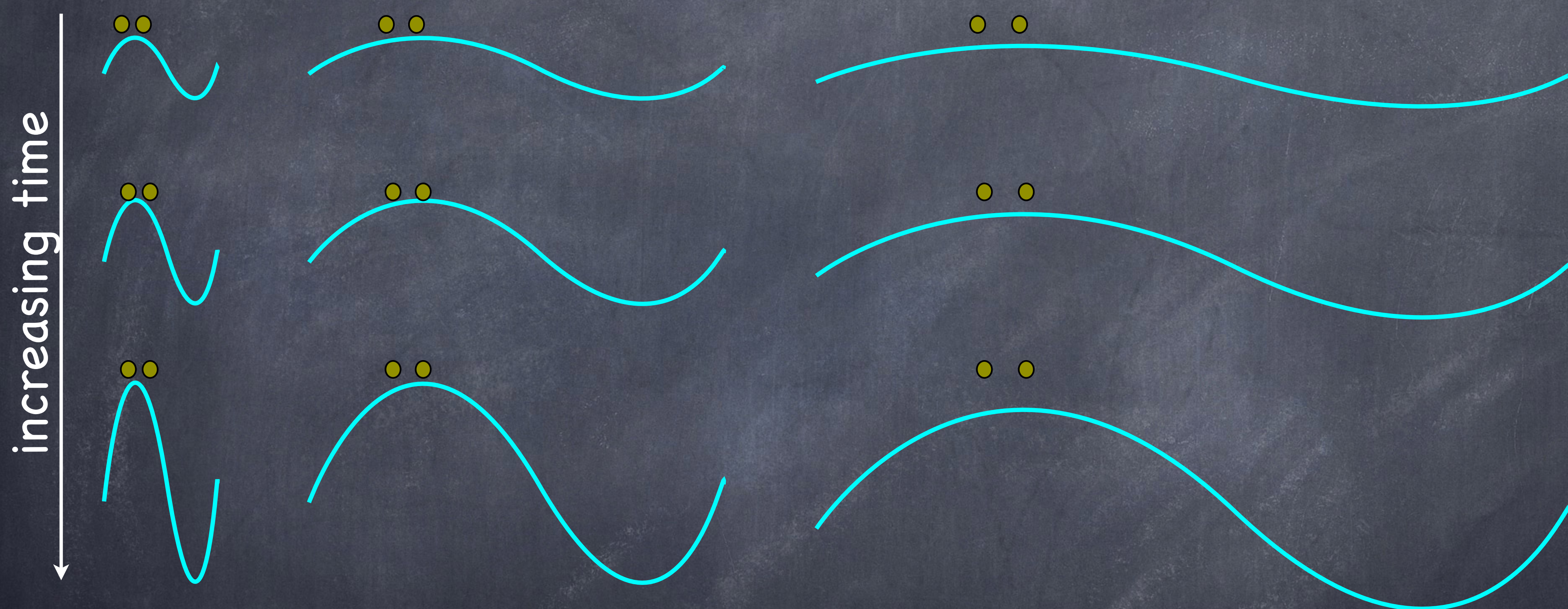
In a universe with CDM only, the linear evolution of matter fluctuations is independent of their wavelength



$$\frac{\delta\rho}{\rho}(k, z_{final}) \propto D(z_{final}) \frac{\delta\rho}{\rho}(k, z_{initial})$$

Main idea:

In a universe with CDM only, the linear evolution of matter fluctuations is independent of their wavelength

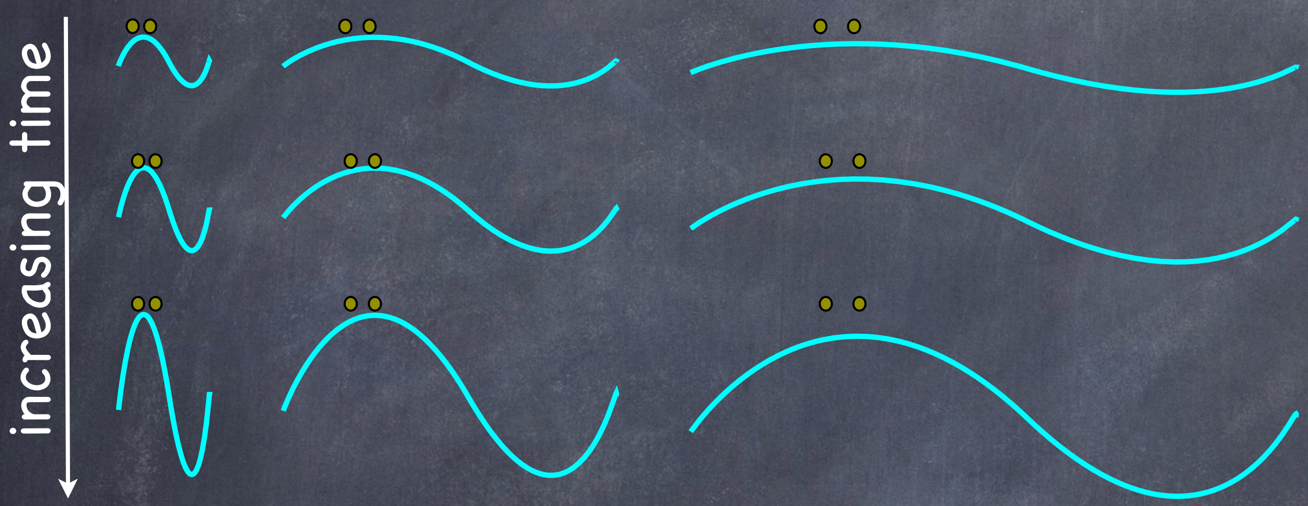


$$\frac{\delta\rho}{\rho}(k, z_{final}) \propto D(z_{final}) \frac{\delta\rho}{\rho}(k, z_{initial})$$

halos can't tell the wavelength of the background matter density perturbation

Main idea:

In a universe with CDM only, the linear evolution of matter fluctuations is independent of their wavelength



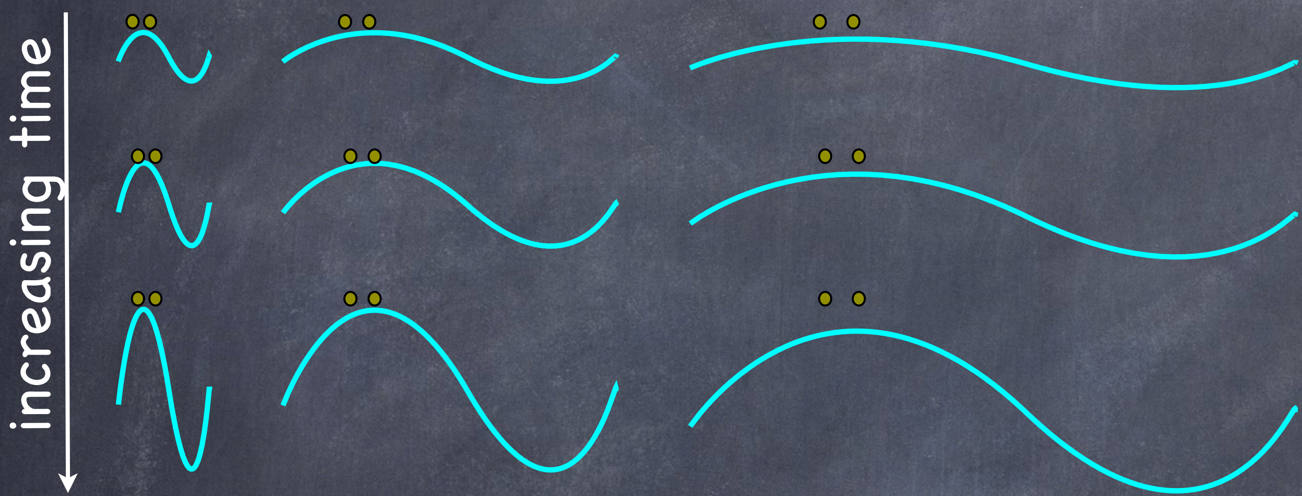
halos can't tell the wavelength of the background matter density perturbation

↓

the effect of $\frac{\delta\rho}{\rho}$ on the halo field (the linear bias) is independent of k

Main idea:

In a universe with CDM only, the linear evolution of matter fluctuations is independent of their wavelength



halos can't tell the wavelength of the background matter density perturbation

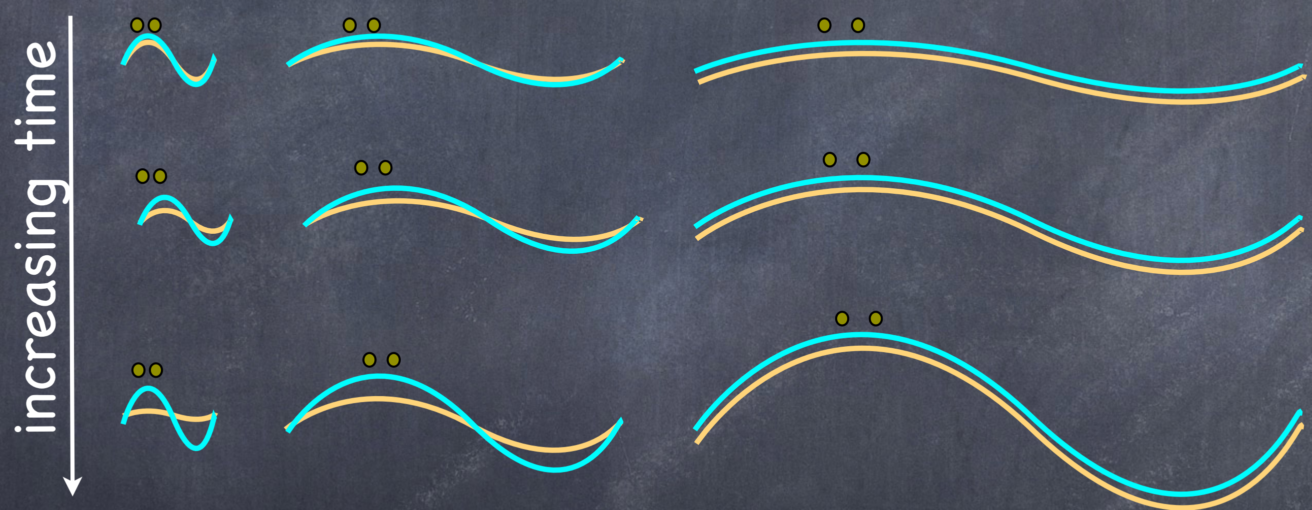
↓

the effect of $\frac{\delta\rho}{\rho}$ on the halo field (the linear bias) is independent of k

massive neutrinos break this

halo bias can depend on k

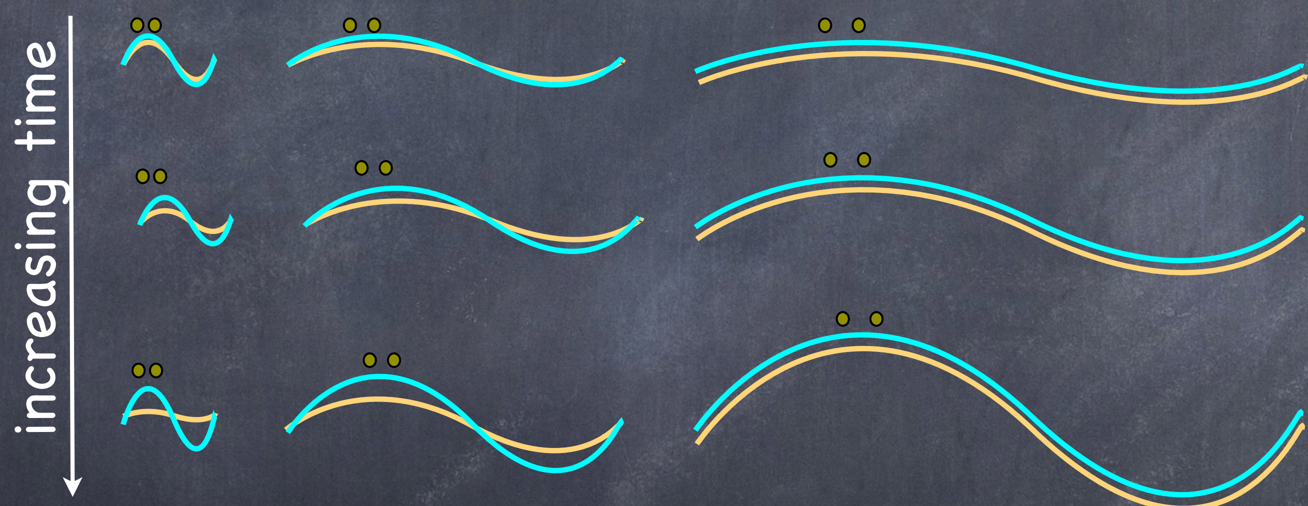
Main idea:



halos
neutrinos
cold dark matter

Main idea:

WANT: estimate of k-dependence of the halo bias caused by massive neutrinos



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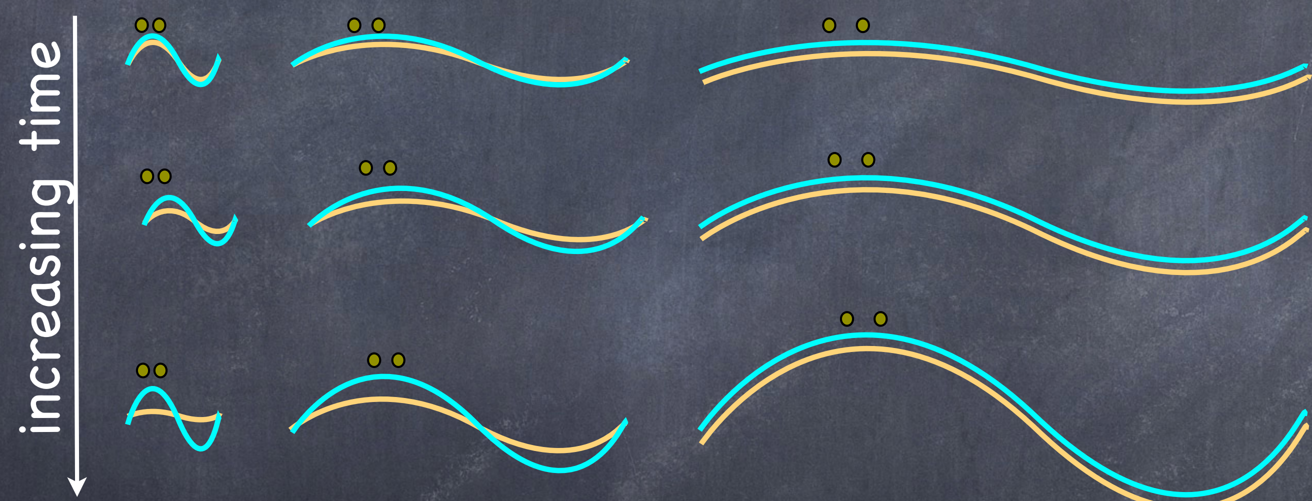


halos
neutrinos
cold dark matter

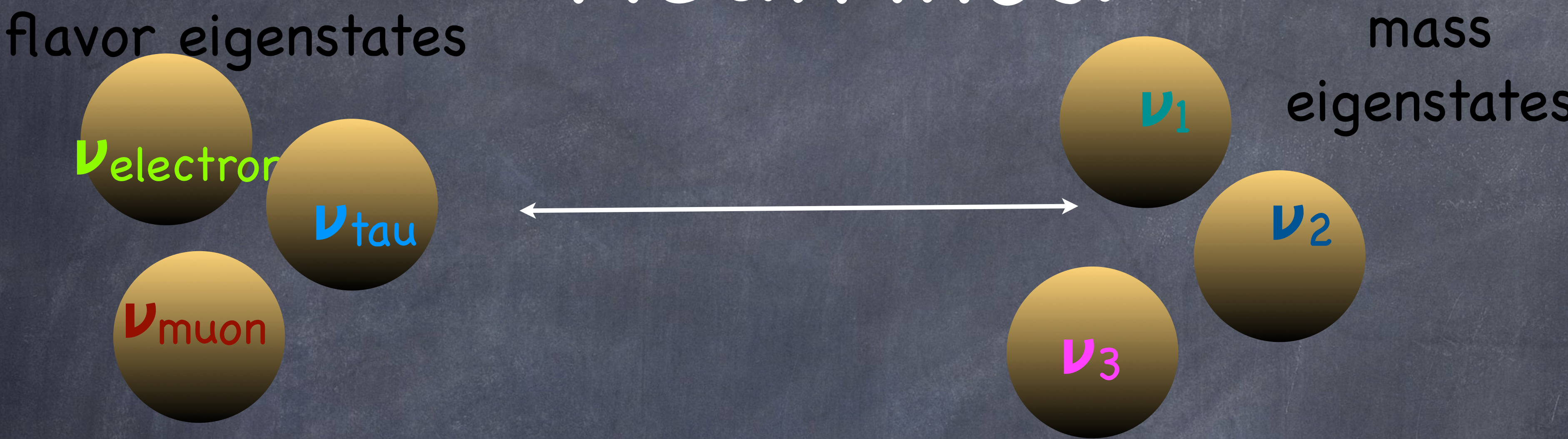
(see also Hui & Parfrey 2008; Parfrey, Hui, Sheth 2011;)

Outline

- Neutrinos in Cosmology
- Scale-dependent structure growth from massive neutrinos
- Scale-dependent halo bias from massive neutrinos
- Observational Consequences

Neutrinos

Neutrinos!



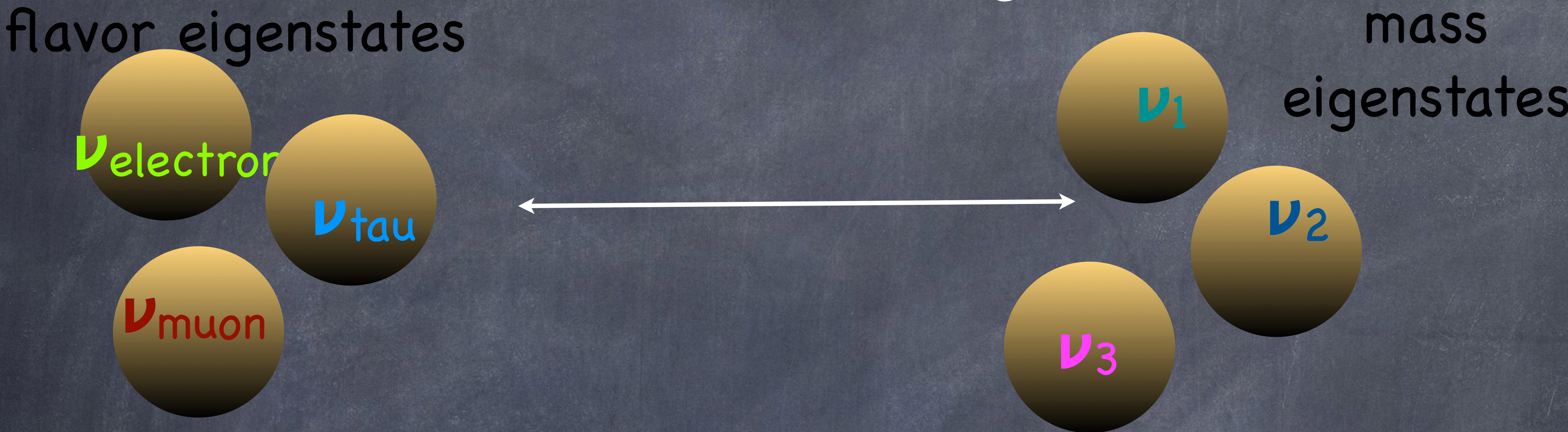
oscillation data gives mass splittings

$m_2^2 - m_1^2 = (7.5 \pm 0.2)\ 10^{-5}\ \mathrm{eV}^2$ (solar neutrino oscillations)

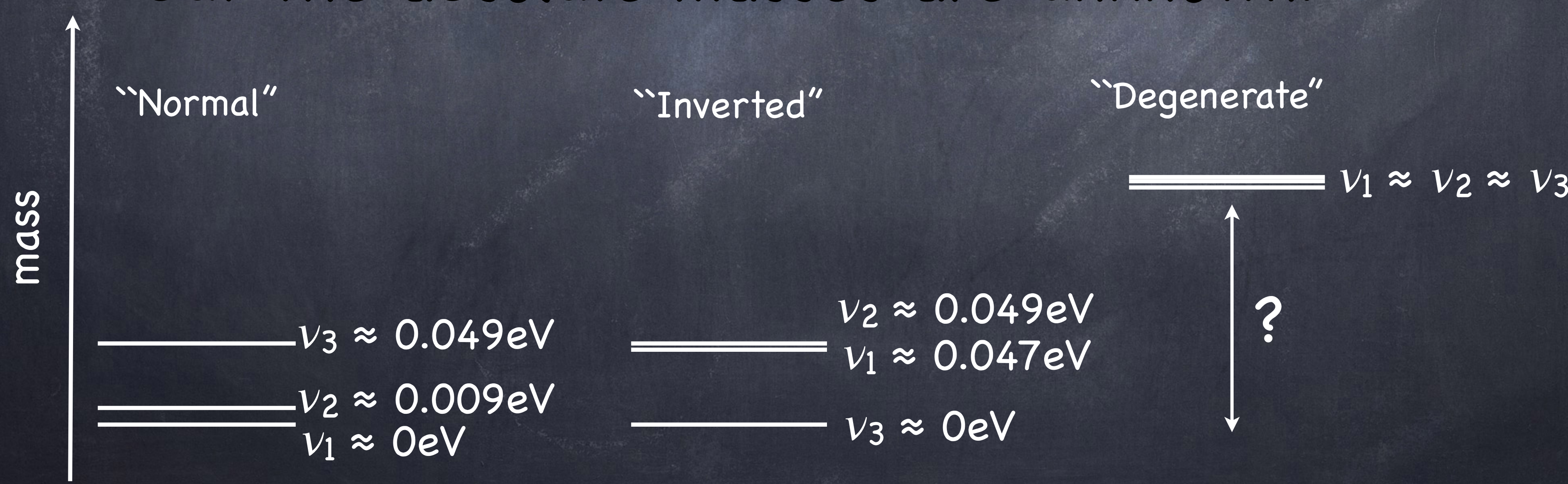
$|m_3^2 - m_2^2| = (2.32^{+0.12}_{-0.08})\ 10^{-3}\ \mathrm{eV}^2$ (atmospheric neutrino oscillations)

Pontecorvo 1957, 1958, 1967; Maki, Nakagawa, Sakata 1962

Neutrinos!



but the absolute masses are unknown!



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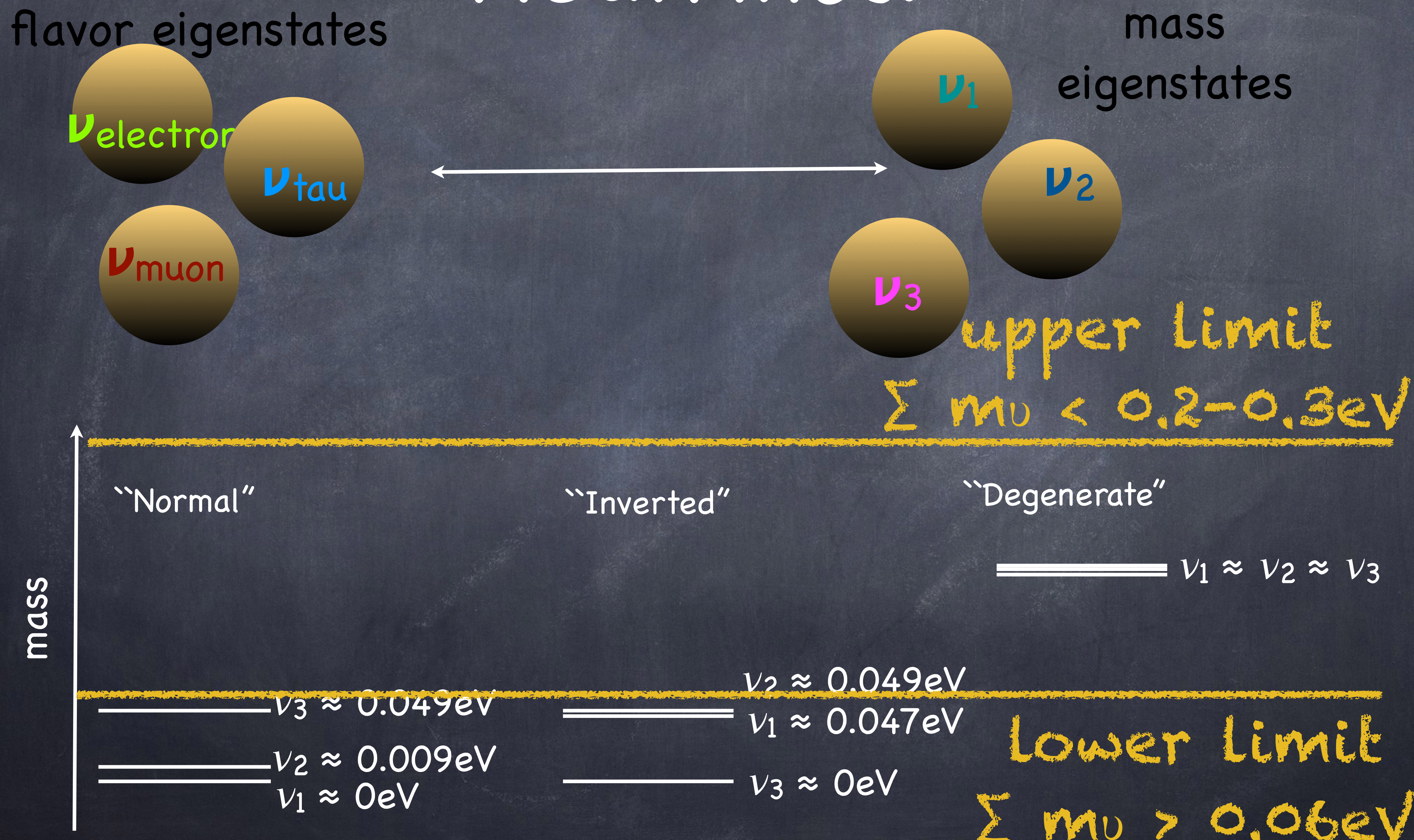

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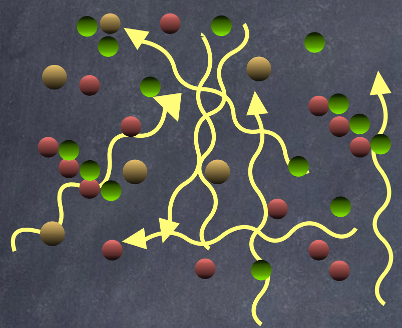
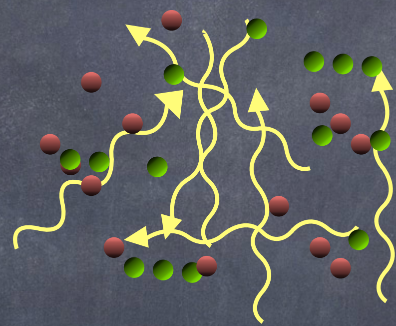
Neutrinos in Cosmology

Neutrinos in cosmology



neutrinos in equilibrium with photons e+, e-

$T_\gamma = T_\nu \propto 1/a$



neutrinos decoupled

$T_\nu \propto 1/a$

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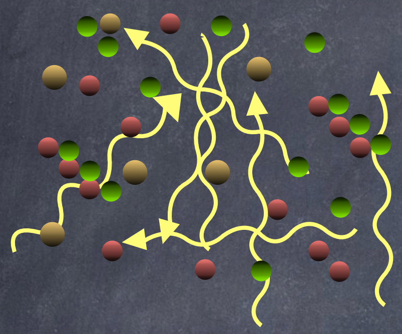
relativistic, in thermal equilibrium at early times

$$n_{1\nu} \sim T_\nu^3$$

$$T_\gamma \approx \left(\frac{11}{4}\right)^{1/3} T_\nu$$

Neutrinos in cosmology



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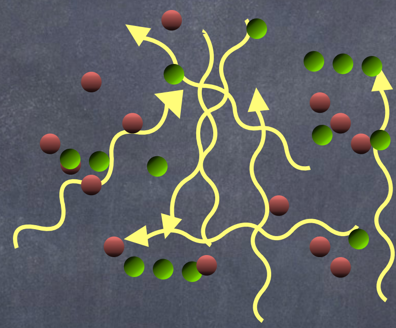

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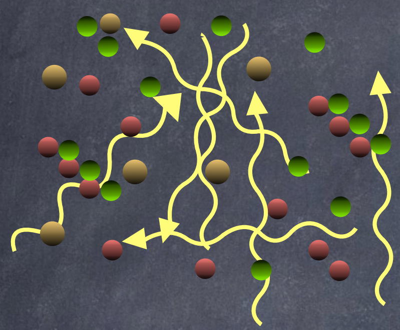
energy density dominated by mass at late times

$$\rho_\nu \sim \sum_i m_{\nu_i} n_{1\nu}$$

Neutrinos in cosmology



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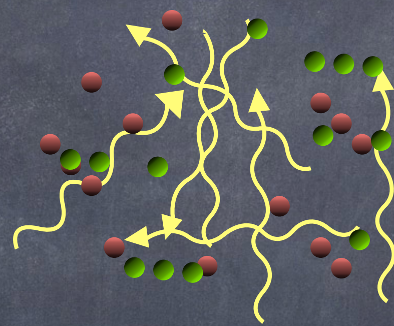

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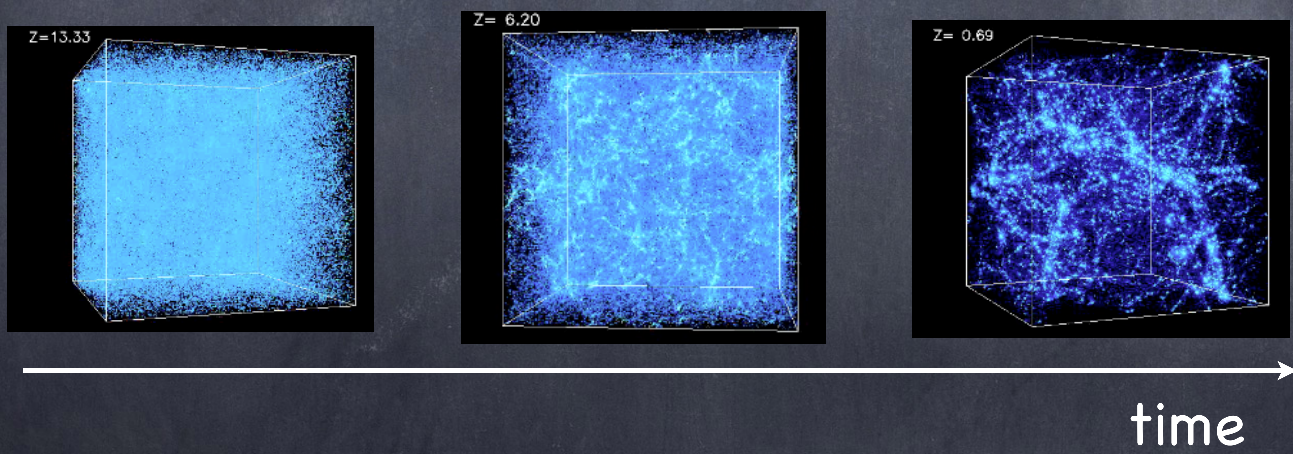
energy density dominated by mass at late times

$$\rho_\nu \sim \sum_i m_{\nu_i} n_{1\nu}$$

$n_{1\nu}$ is known, so a measurement of ρ_ν gives $\sum m_\nu$

Neutrinos in Large-scale structure



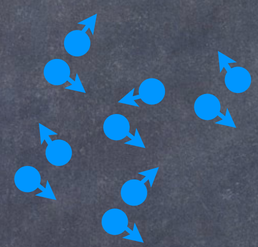
(Kravtsov)

Massive neutrinos and linear structure growth

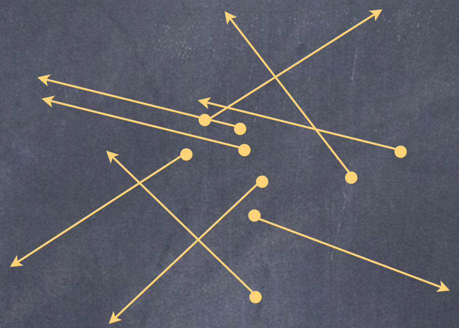
The gravitational evolution of large-scale structure is different for fast and slow moving particles

Massive neutrinos and linear structure growth

The gravitational evolution of large-scale structure is different for fast and slow moving particles



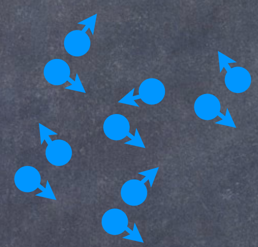
(clump easily)



(don't clump easily)

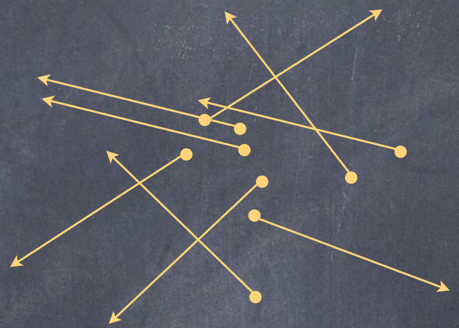
Massive neutrinos and linear structure growth

The gravitational evolution of large-scale structure is different for fast and slow moving particles



(clump easily)

baryons and cold dark matter



(don't clump easily)

neutrinos (or other exotic light dark matter)

Massive neutrinos and linear structure growth

small-scale density perturbations don't retain neutrinos

cold dark matter and baryons density perturbation growing

neutrino density perturbation decaying

time

$\frac{\delta\rho_c}{\rho_c}$

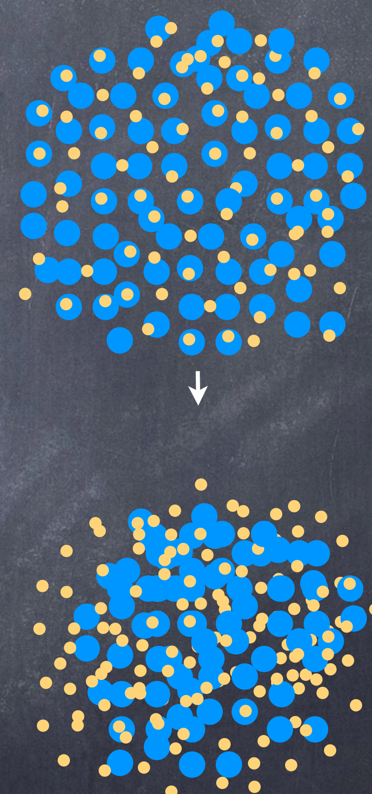
$\frac{\delta\rho_\nu}{\rho_\nu}$

Massive neutrinos and linear structure growth

small-scale density perturbations don't retain neutrinos

time

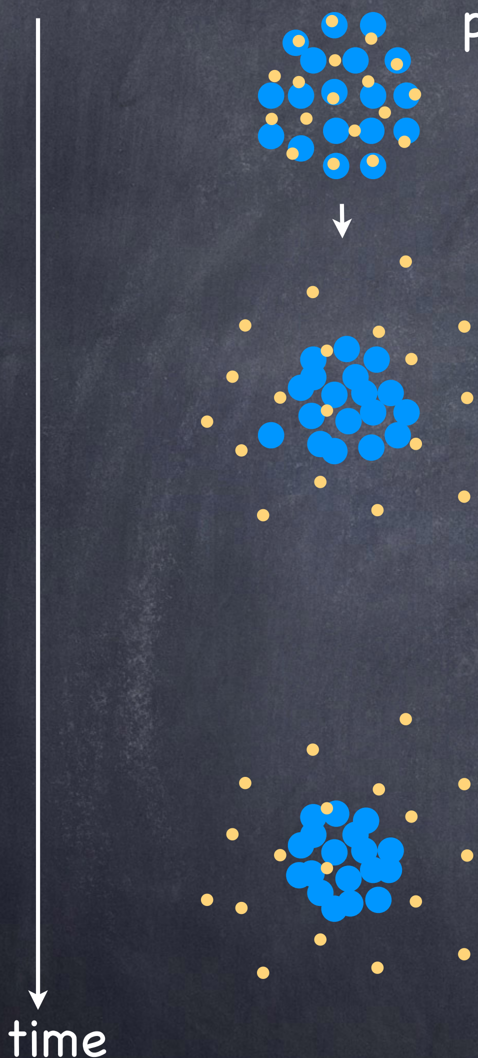
large-scale density perturbations do retain neutrinos



cold dark matter, baryons and neutrinos growing together

$$\frac{\delta\rho_\nu}{\rho_\nu} \quad \frac{\delta\rho_c}{\rho_c}$$

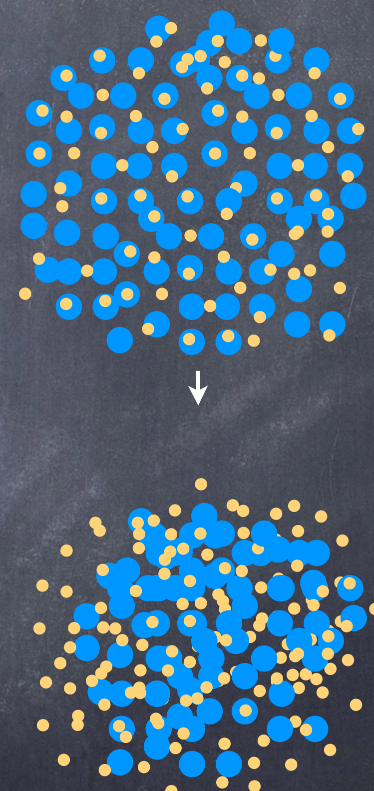
Massive neutrinos and linear structure growth



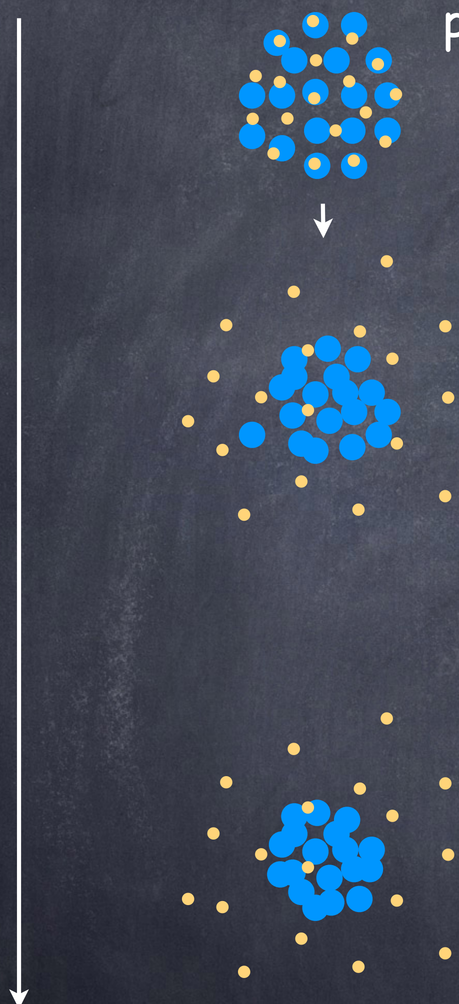
small-scale density perturbations don't retain neutrinos

Growth of matter perturbations is scale-dependent

large-scale density perturbations do retain neutrinos



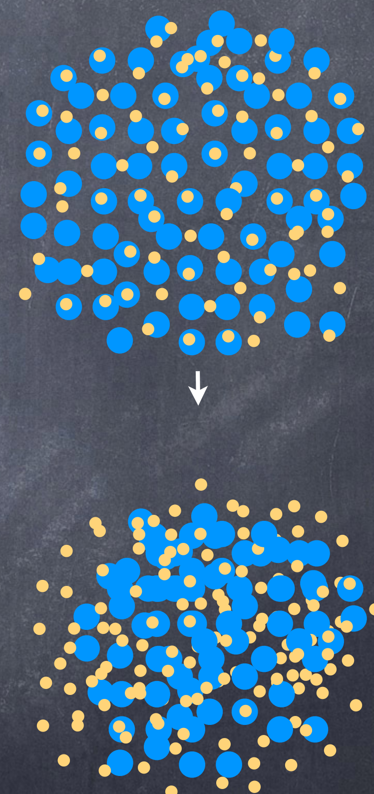
Massive neutrinos and linear structure growth



time

small-scale density perturbations don't retain neutrinos

large-scale density perturbations do retain neutrinos



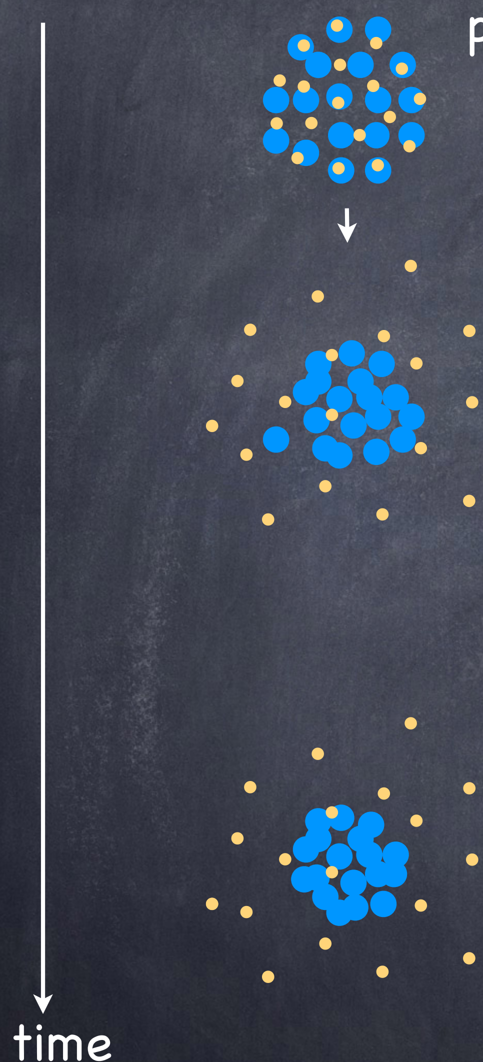
Growth of matter perturbations is scale-dependent

Relevant scale:

Typical distance a neutrino can travel in a Hubble time

$\lambda_{fs} \sim u_\nu/H$

Massive neutrinos and linear structure growth



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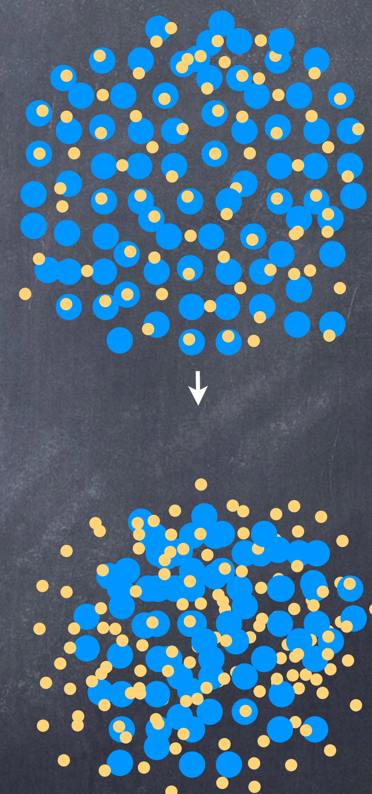


Growth of matter perturbations is scale-dependent

Relevant scale:

Typical distance a neutrino can travel in a Hubble time

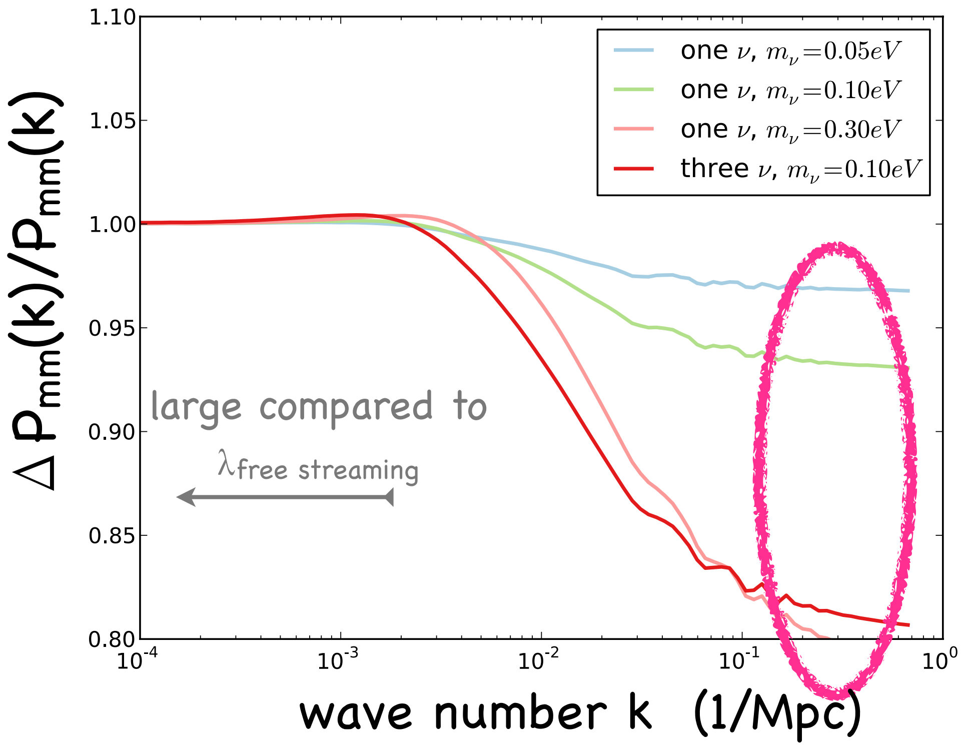
$\lambda_{fs} \sim u_\nu / H$

↑

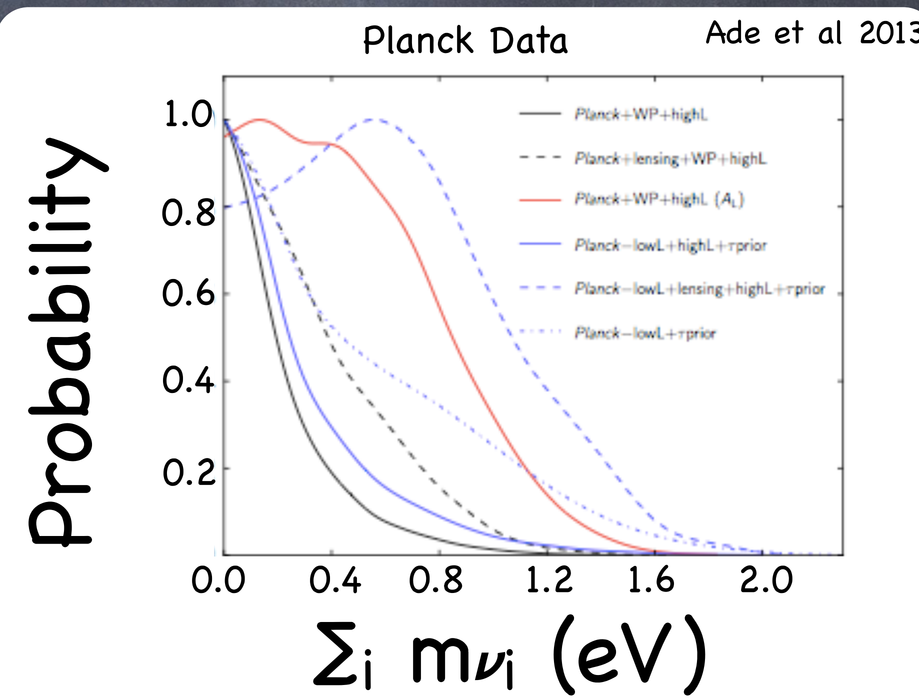
"free-streaming scale"

Scale-dependent growth

change in typical amplitude of $\delta_m(k)$ from $m_\nu \neq 0$



cosmological constraints!



Bond, Efstathiou, Silk 1980
Hu, Eisenstein, Tegmark 1998

The scale-dependent growth of density perturbations causes halo bias to be scale dependent

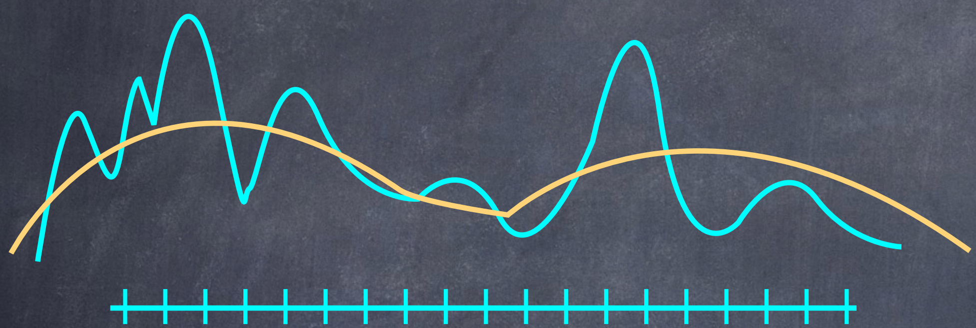
The scale-dependent growth of density perturbations causes halo bias to be scale dependent

$$\frac{\delta n}{n}(k) = b\,\frac{\delta\rho_m}{\rho_m}(k) \longrightarrow \frac{\delta n}{n}(k) = b(k)\,\frac{\delta\rho_m}{\rho_m}(k)$$

Prescription for calculating the halo bias in a universe with massive neutrinos

Prescription for calculating the halo bias

initial density field

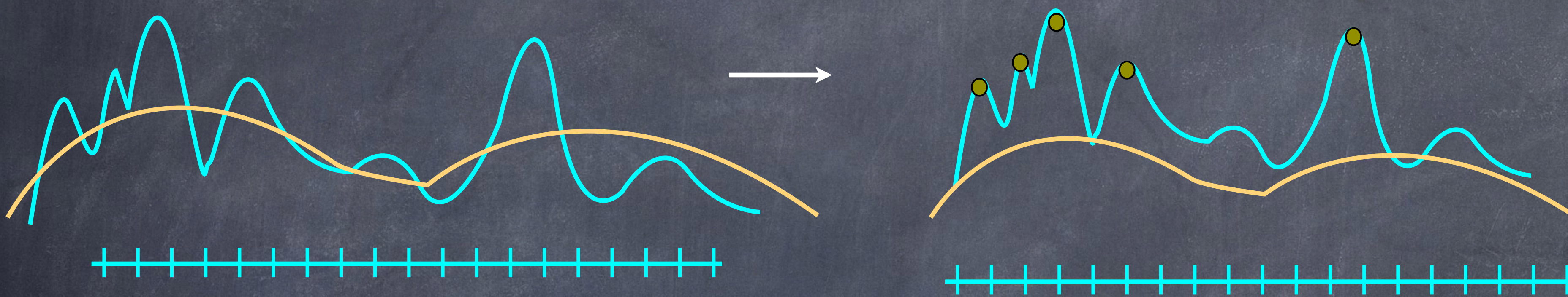


(ML 2014)

Gunn & Gott 1972
Press & Schechter 1974

Prescription for calculating the halo bias

initial density field　　　　initial proto-halo distribution



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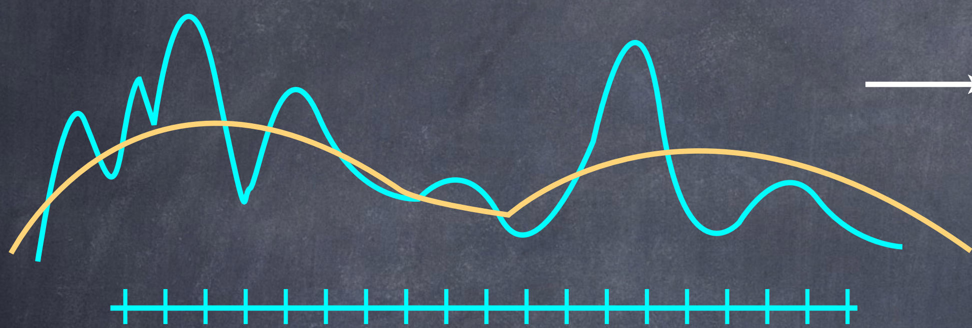
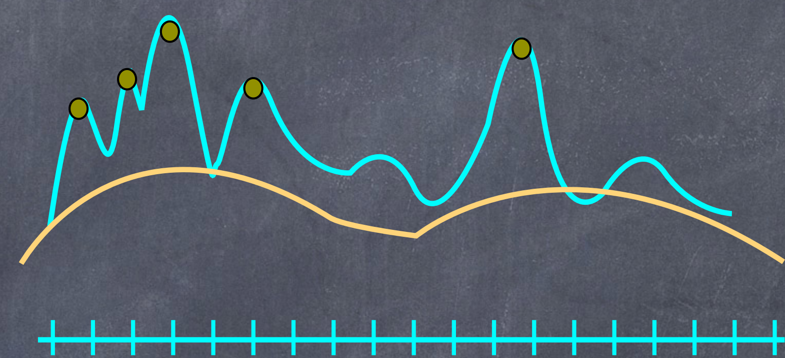
Prescription for calculating the halo bias

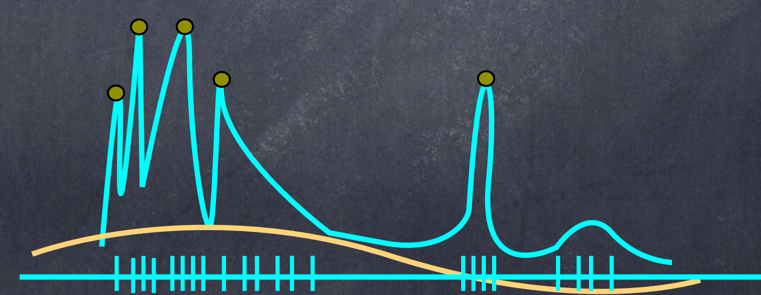
initial density field



initial proto-halo distribution



late time halo distribution

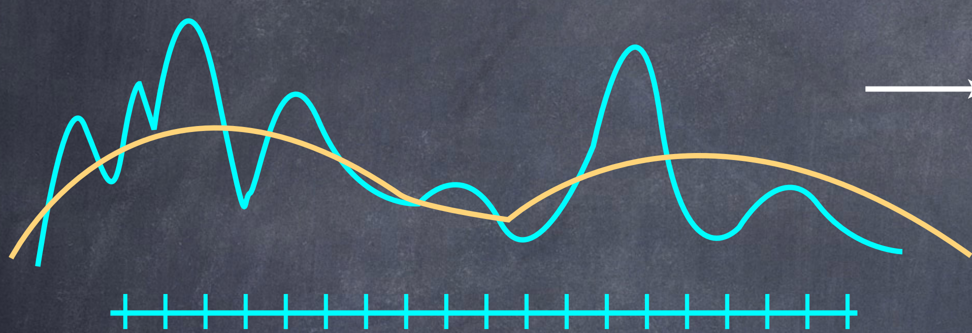


(ML 2014)

Gunn & Gott 1972
Press & Schechter 1974

Prescription for calculating the halo bias

initial density field



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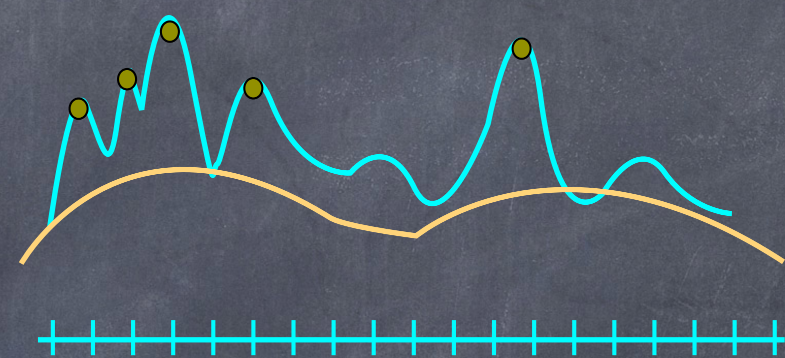
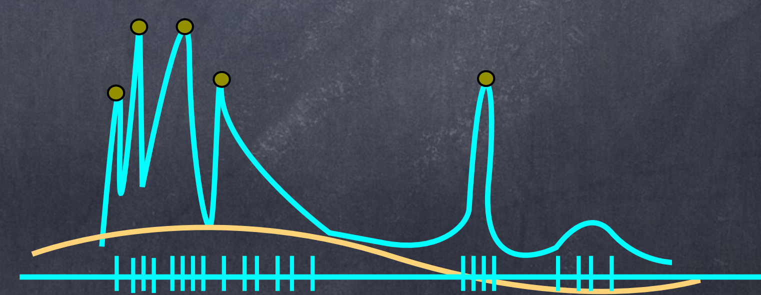


late time halo distribution



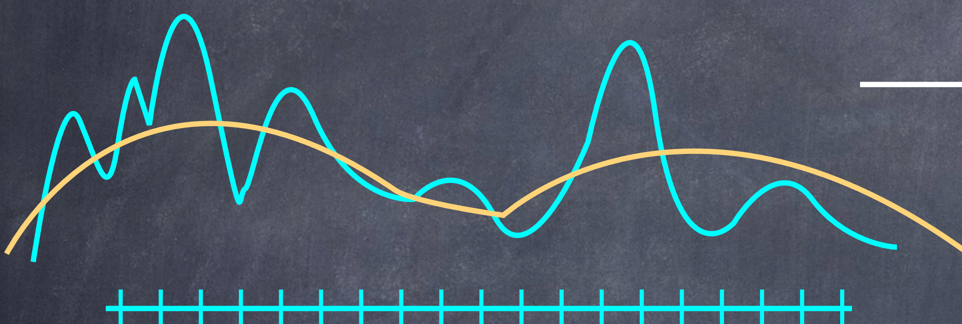
$$\frac{\delta n}{n} \equiv b \left.\frac{\delta\rho}{\rho}\right|_{\text{long-wavelength}}$$

(ML 2014)

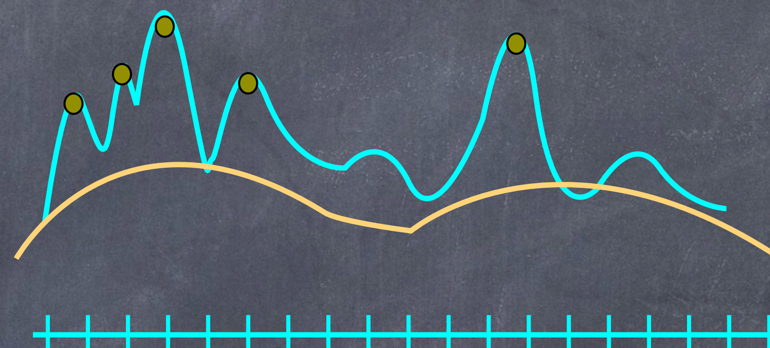
Gunn & Gott 1972
Press & Schechter 1974

Prescription for calculating the halo bias

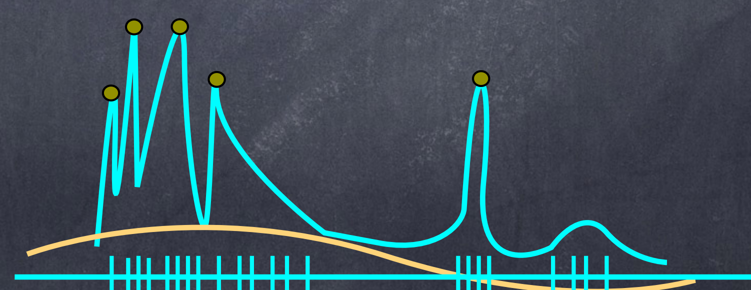
initial density field



initial proto-halo distribution



late time halo distribution



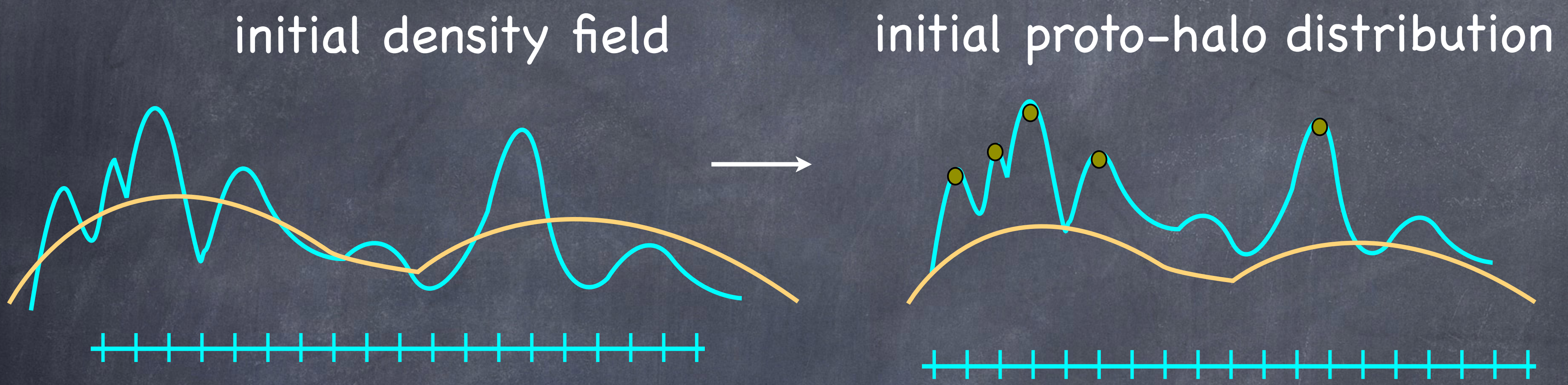
$$\frac{\delta n}{n} \equiv b \left.\frac{\delta\rho}{\rho}\right|_{\text{long-wavelength}}$$

want this!

(ML 2014)

Gunn & Gott 1972
Press & Schechter 1974

Prescription for calculating the halo bias



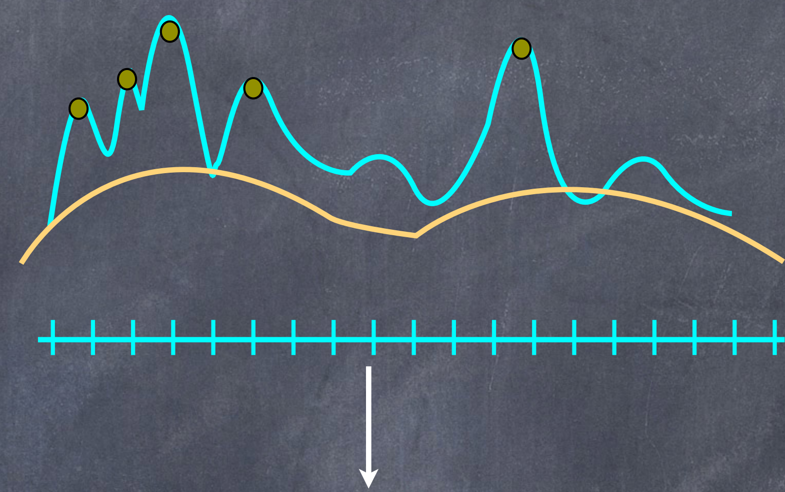
(i) In this step, the scale-dependent evolution of density perturbations causes

$$\frac{\delta n_{Lagrangian}}{n}(k) = b_{Lagrangian}(k)\,\frac{\delta\rho_{cdm}}{\rho_{cdm}}(k)$$

(ML 2014)

Prescription for calculating the halo bias

initial proto-halo distribution

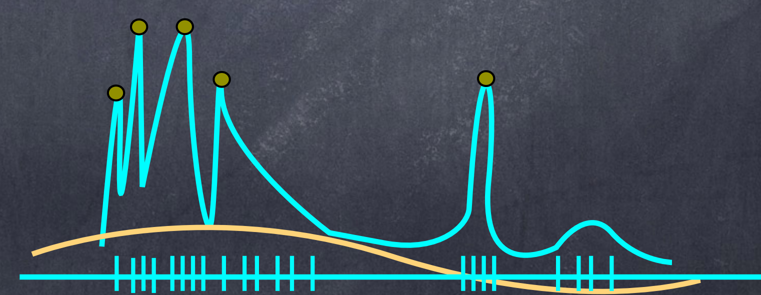


(ii) In this step, the free streaming of neutrinos causes

$$b(k) = \frac{\langle \delta_n(k)\delta_m(k)\rangle}{\langle \delta_m(k)\delta_m(k)\rangle}$$

$$= (1 + b_{Lagrangian}(k))\frac{\langle \delta_{cdm}(k)\delta_m(k)\rangle}{\langle \delta_m(k)\delta_m(k)\rangle}$$

late time halo distribution



(ML 2014)

Prescription for calculating the halo bias

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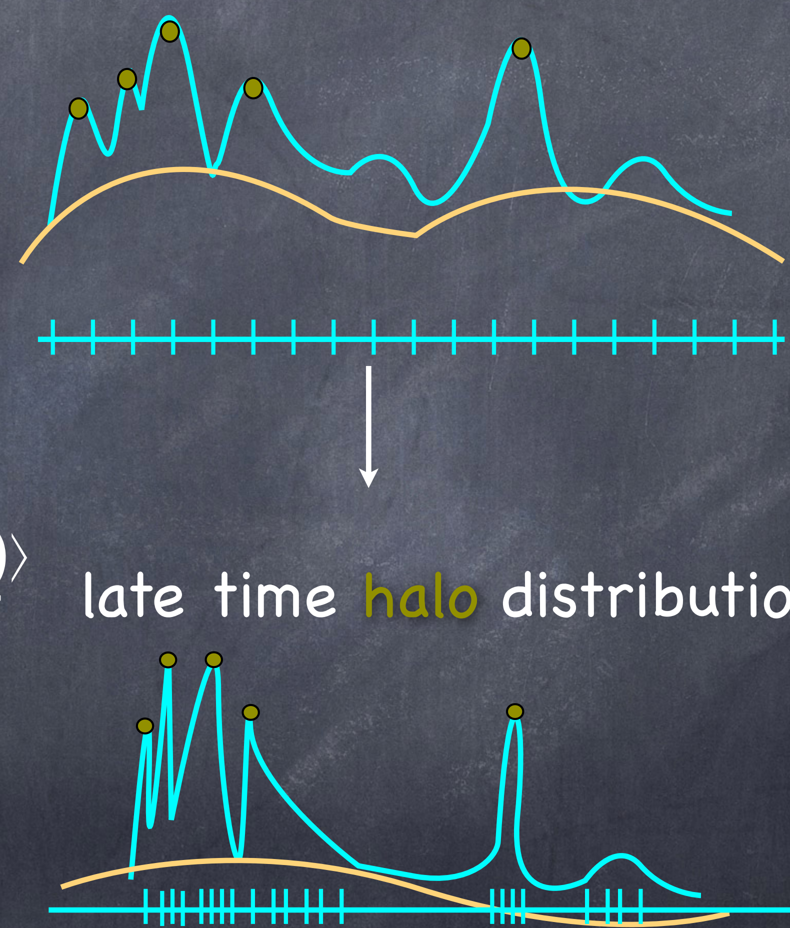
(i.e. halos trace CDM —> bias w.r.t total matter is scale-dependent)

(ML 2014)

Prescription for calculating the halo bias

initial proto-halo distribution



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$$b(k) = \frac{\langle \delta_n(k)\delta_m(k)\rangle}{\langle \delta_m(k)\delta_m(k)\rangle}$$

$$= (1 + b_{\mathrm{Lagrangian}}(k))\frac{\langle \delta_{cdm}(k)\delta_m(k)\rangle}{\langle \delta_m(k)\delta_m(k)\rangle}$$

(i.e. halos trace CDM —> bias w.r.t total matter is scale-dependent)

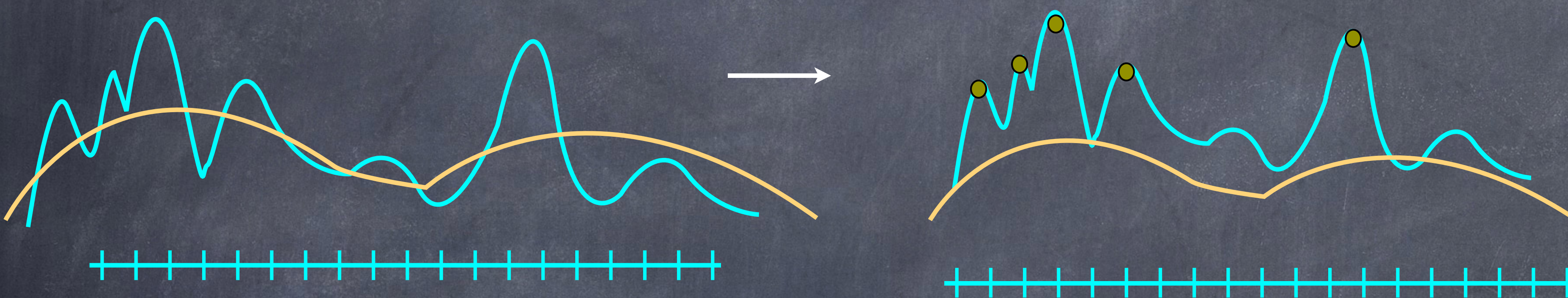
see also Villaescusa-Navarro, Marulli, Viel, Branchini, Castorina 2013
Castorina, Sefusatti, Sheth, Villaescusa-Navarrow, Viel 2014
Biagetti, Desjacques, Kehagias, Riotto 2014

(ML 2014)

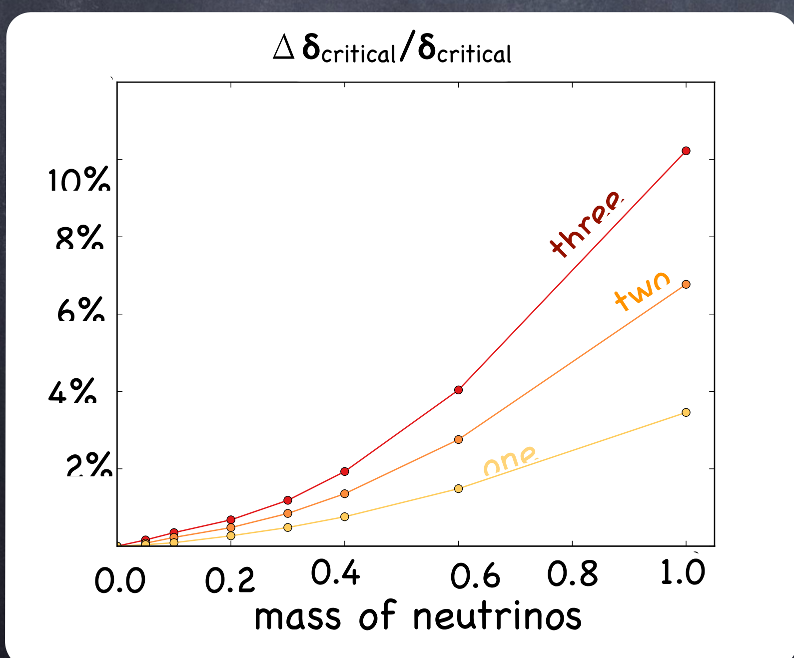
Prescription for calculating the halo bias

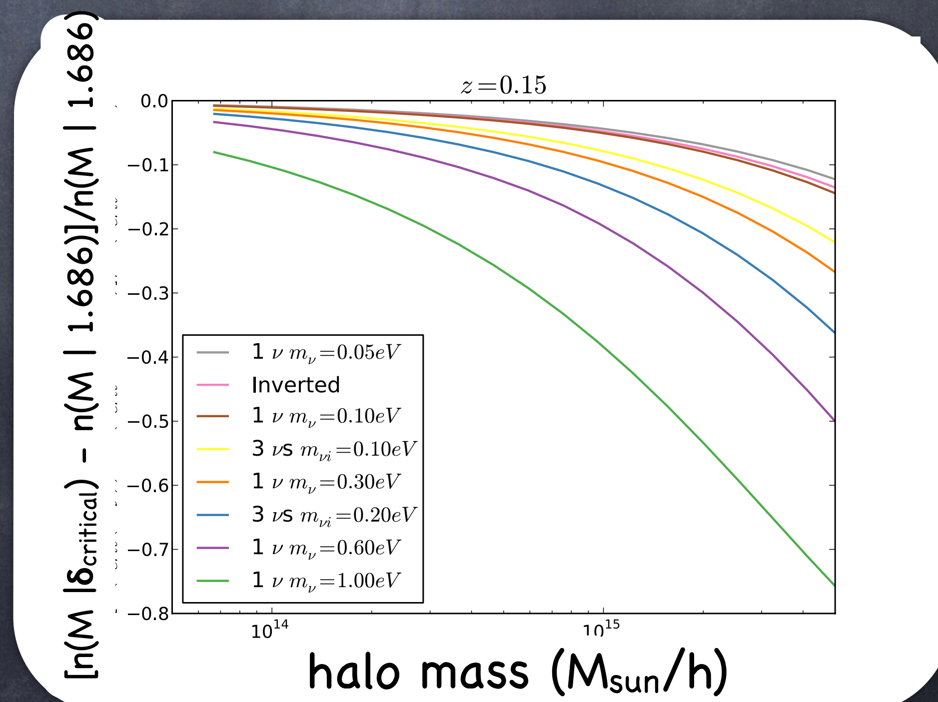
initial density field

initial proto-halo distribution



(this step gives a prediction for n(M): ML 1405:4858)





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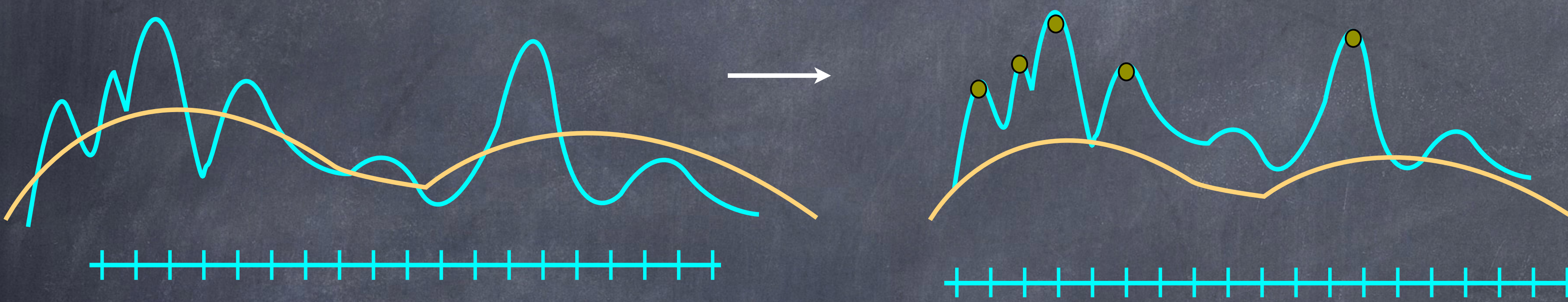
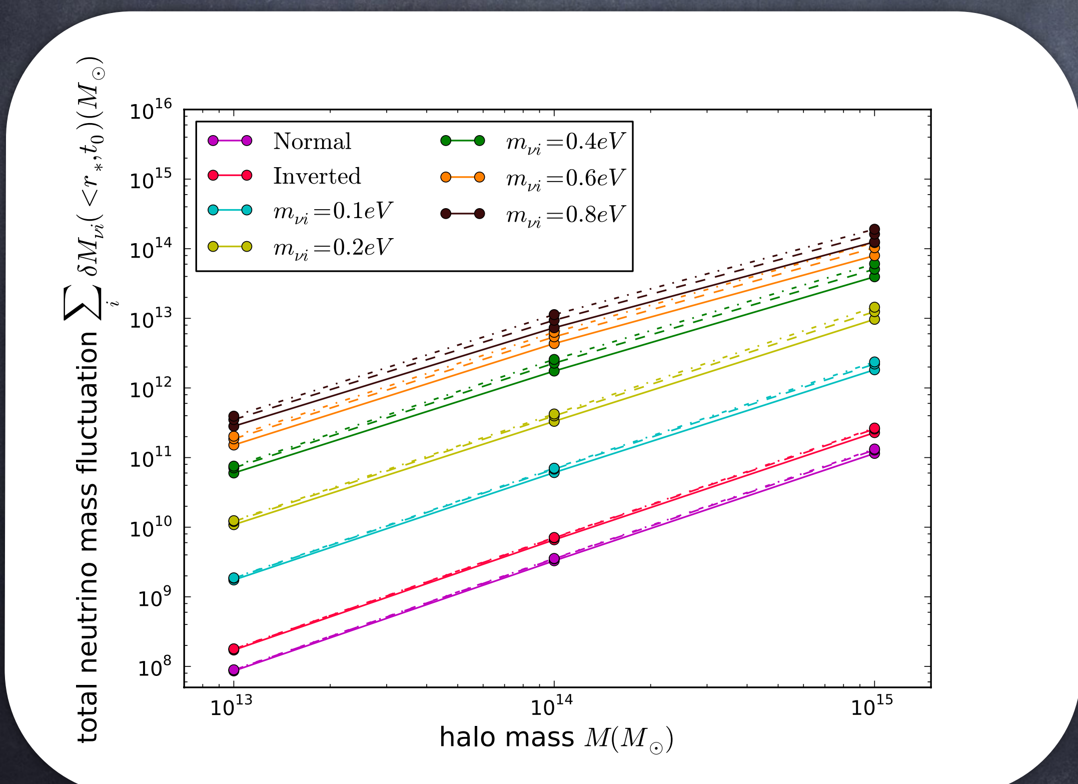




And these are ingredients for calculating the accumulated neutrino mass around CDM halos

ML & Zaldarriaga 1310:6459

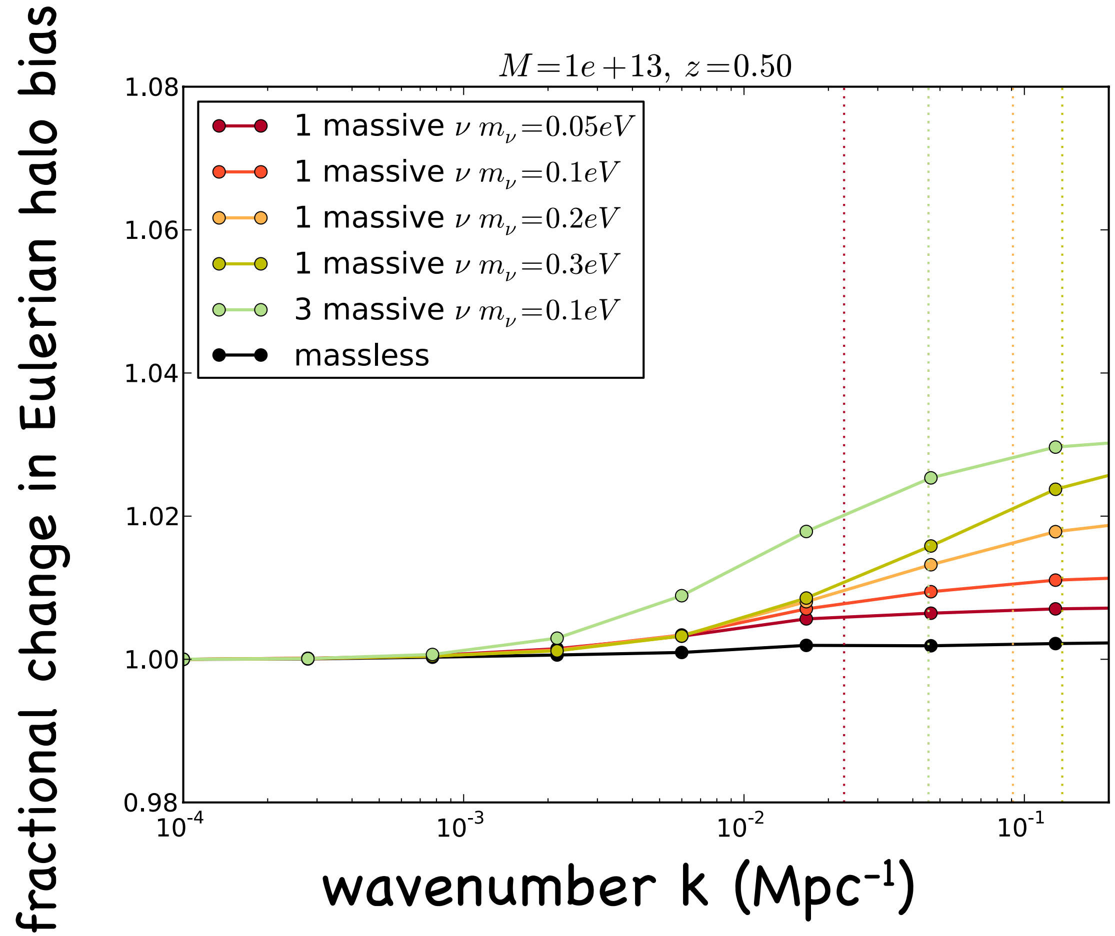
Numerical estimates for scale-dependent halo bias

Numerical results for halo bias

scale-dependent change to final bias

$$\delta n(k)/n = b(k)\ \delta_{matter}(k)$$

$$b(k) = \sqrt{P_{hh}(k)/P_{mm}(k)}$$



(Use Bhattacharya et al 2011 for $n(M|\ \delta_{crit})$)

(ML 2014)

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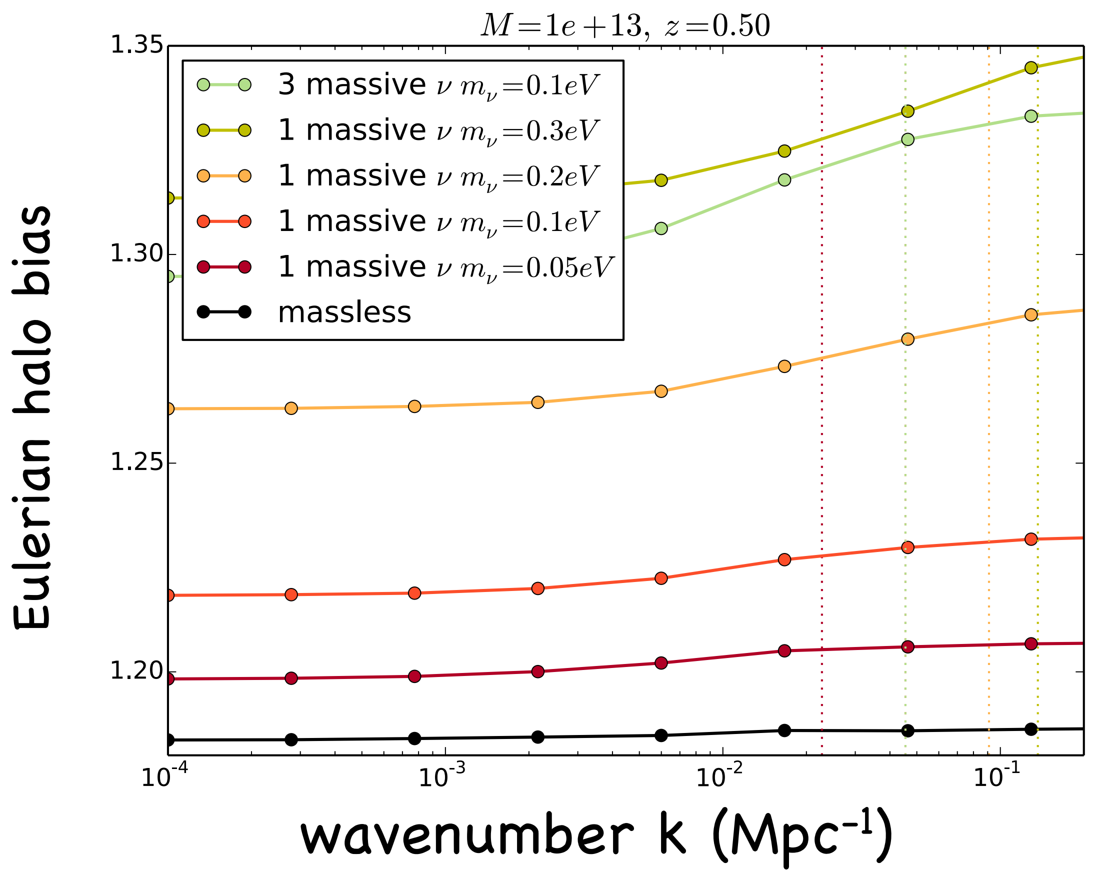

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(ML 2014)

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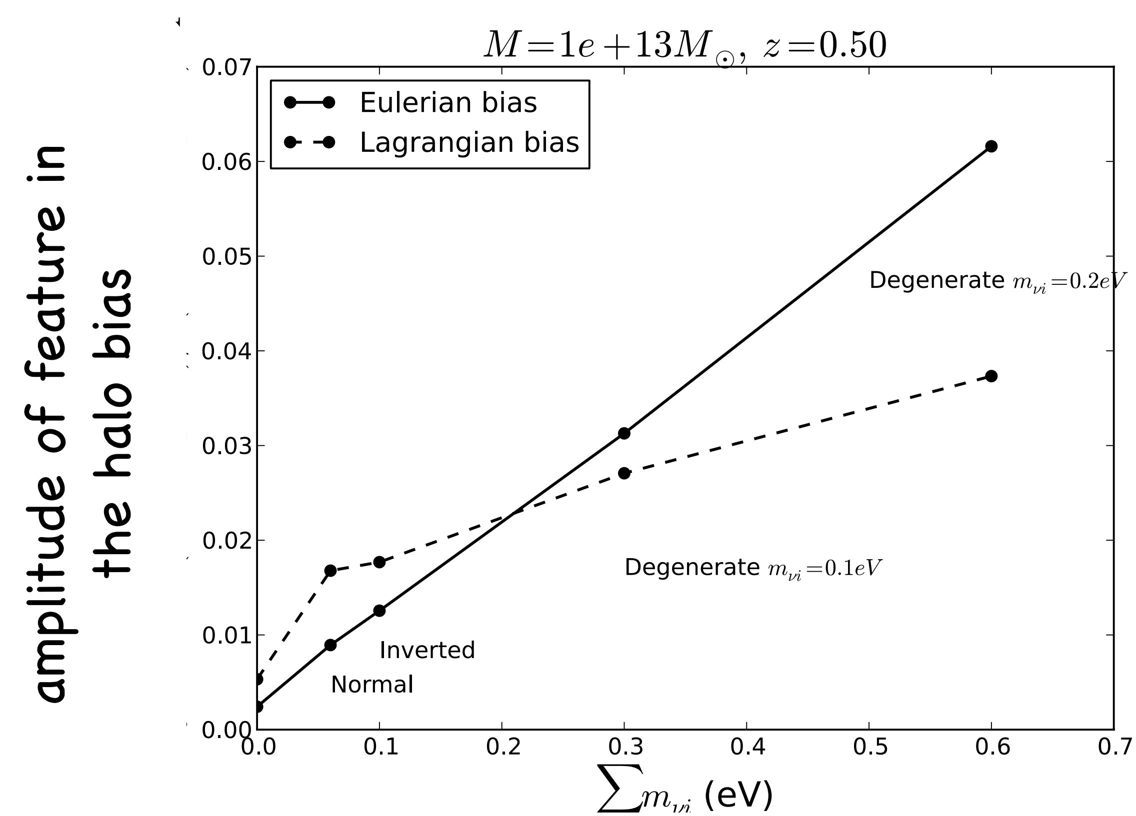

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(ML 2014)

Observational consequences of scale-dependent bias?

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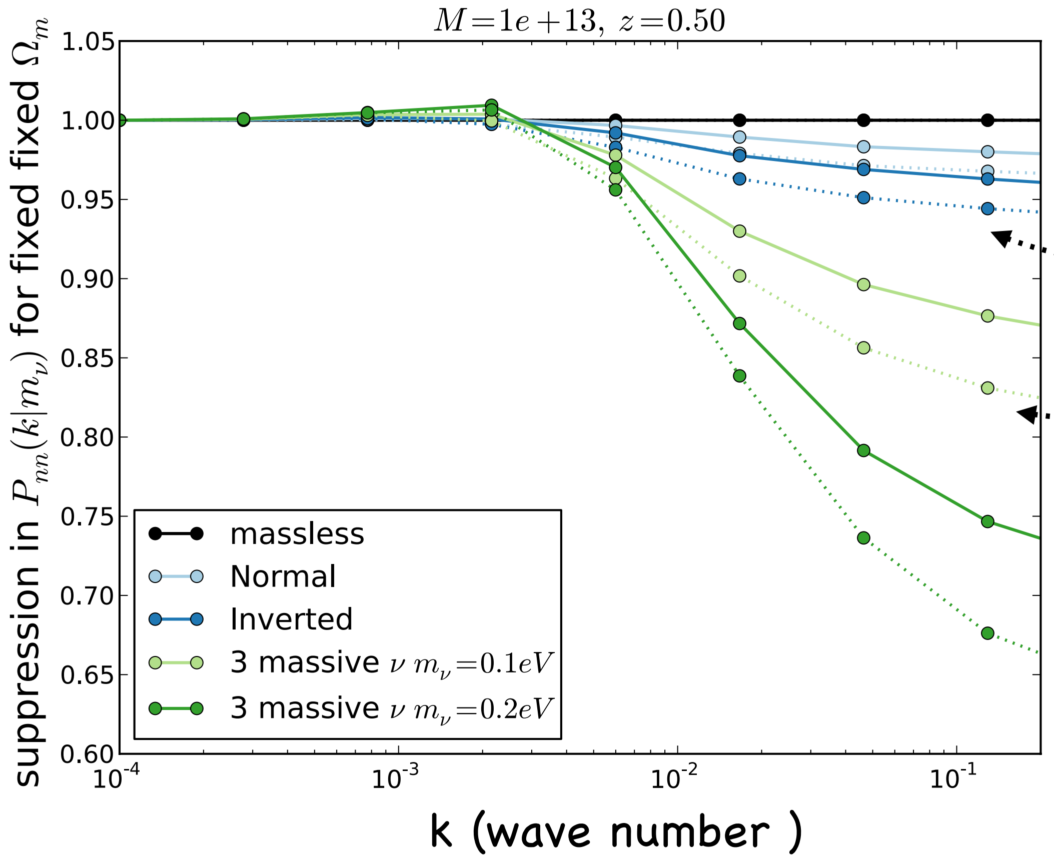
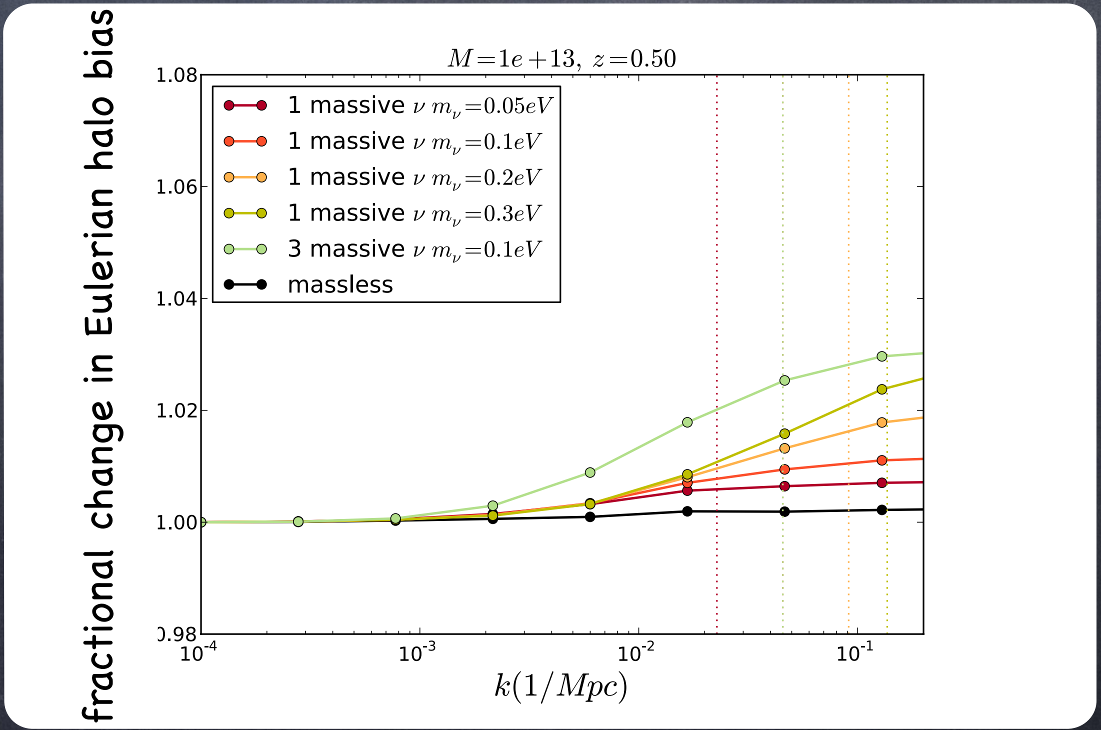
(incorrectly) assuming constant bias

suppression in galaxy power spectrum less than in matter power spectrum

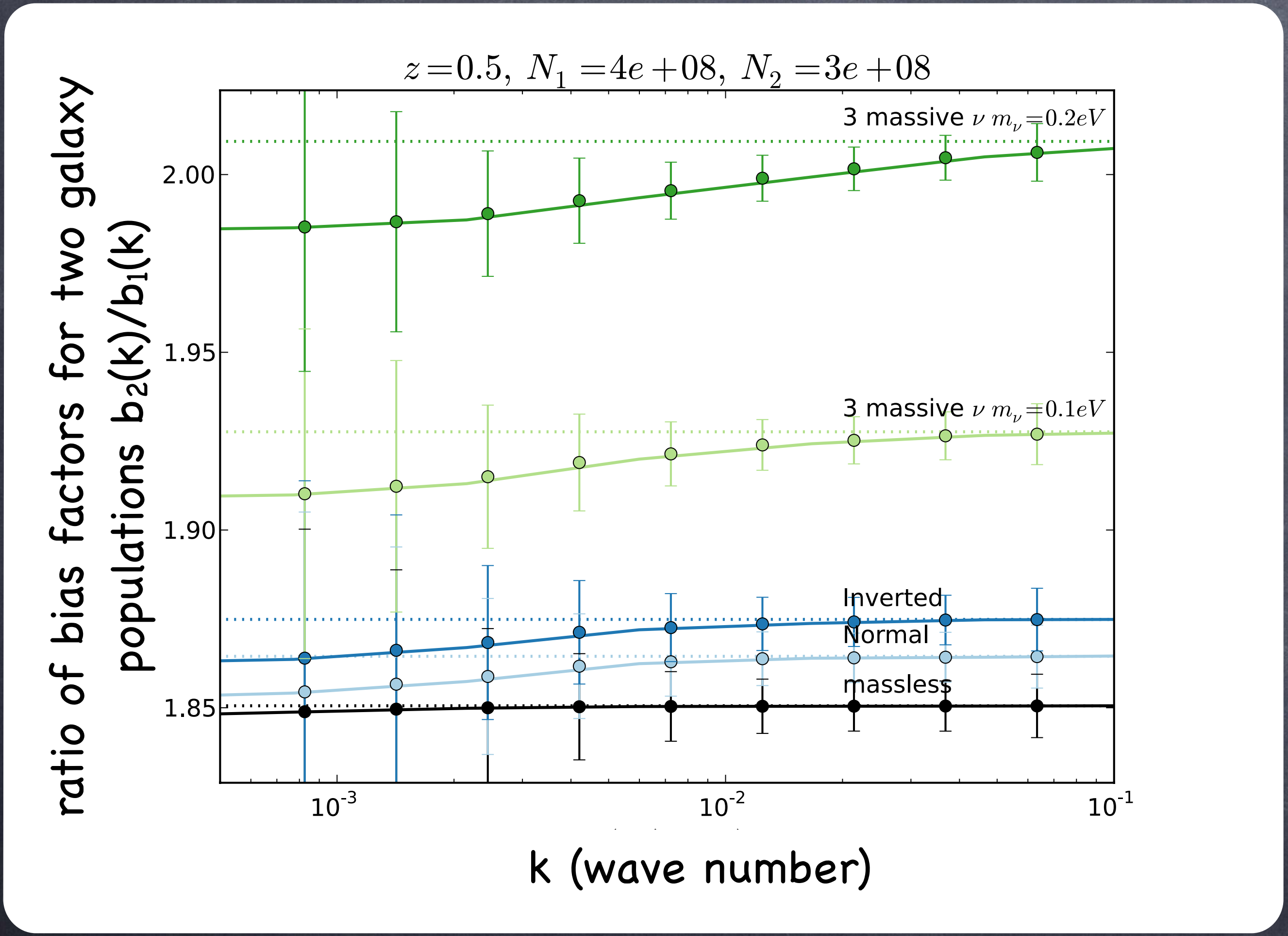


(ML 2014)

But the scale-dependent halo bias is itself an observable!



The scale-dependent halo bias is an observable!



$$\sigma_{b1/b2} \sim \frac{1}{\sqrt{n_1 P_{g2g2}}}$$

(ML in prep.)

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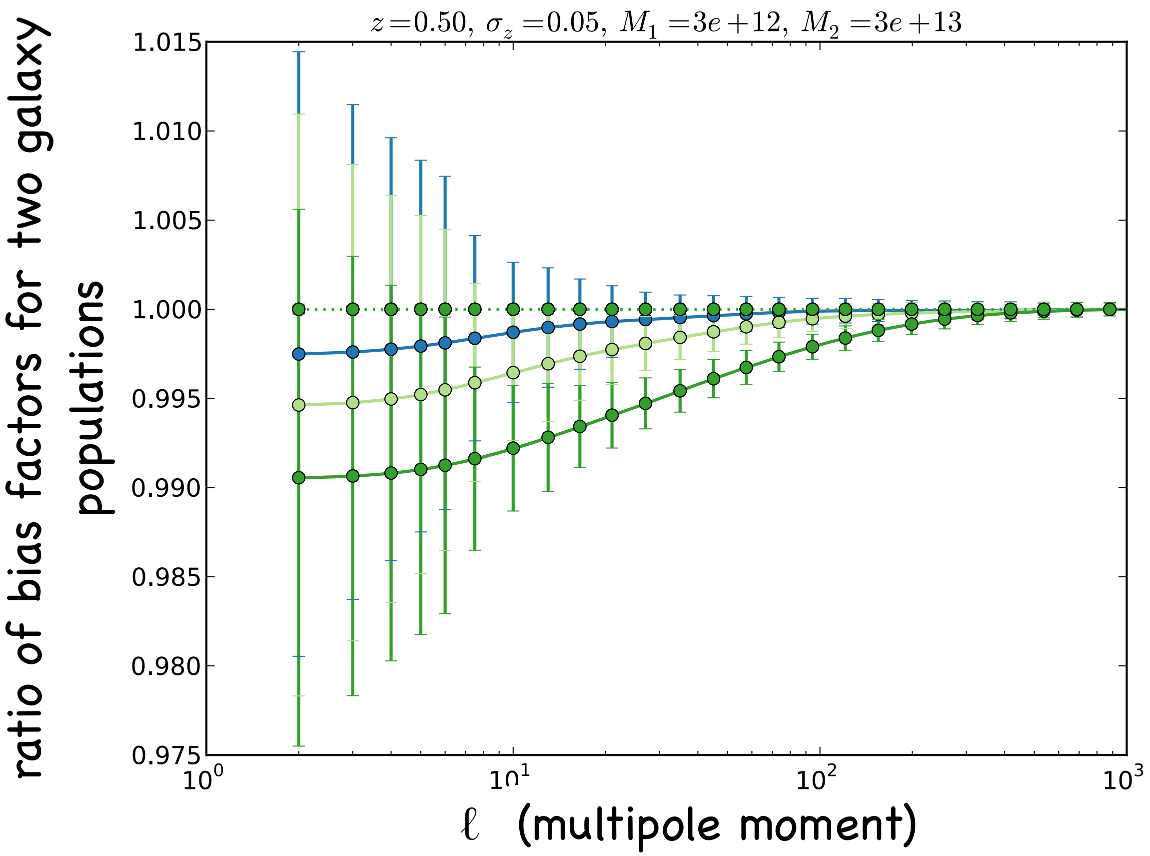

$$\sigma_{b1/b2} \sim \frac{1}{\sqrt{N_\ell n_1 C_{g2g2}}}$$

(ML in prep.)

Conclusions

- Cosmology provides interesting information about neutrino physics!
- Scale-dependent halo bias is a new signal of massive neutrinos in large-scale structure
- Scale-dependent halo bias is a new systematic for massive neutrinos in large-scale structure